

Mass-varying massive gravity with k-essence

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Abstract For a large class of mass-varying massive-gravity models, the graviton mass cannot provide the late-time cosmic expansion of the universe due to its vanishing at late time. In this work, we propose a new class of mass-varying massive gravity models, in which the graviton mass varies according to a kinetic term of a k-essence field. By using a more general form of the fiducial metric, we found a solution such that a non-vanishing graviton mass can drive the accelerated expansion of the universe at late time. We also perform dynamical analyses of such a model and find that without introducing the k-essence Lagrangian, the graviton mass can be responsible for both dark contents of the universe, namely dark energy, which drives the accelerated expansion of the universe, and non-relativistic matter, which plays the role of dark matter. Moreover, by including the k-essence Lagrangian, we find that it is possible to alleviate the so-called cosmic coincidence problem.

1 Introduction

Massive gravity has its own series of developments as a modified gravity beyond general relativity. Back in 1939, Fierz and Pauli investigated a first model of massive gravity [1]. The model was a linearized general relativity, where the fluctuation of geometry propagates a spin-2 graviton, plus linear interactions, which, in particle physics language, corresponds to giving a non-zero mass to the graviton; hence the name “massive gravity”. This model was supposed to coincide with general relativity in the massless limit but it faced a theoretical crisis when discontinuities in such a limit were found by van Dam et al. [2,3]. In particular, the discontinuities were found as different predictions between Fierz–Pauli massive gravity and general relativity. The problem remained unsolved for several years, until Vainshtein proposed a way out by introducing higher-order interactions into the Fierz–

Pauli massive gravity [4]. In other words, he claimed that within a particular scale, coined the Vainshtein radius, any predictions from the linear theory cannot be trusted unless nonlinear contributions are taken into account. However, adding such nonlinearities, claimed by Boulware and Deser, not only fixes the discontinuity problem but also introduces a theoretical inconsistency, namely a Boulware–Deser ghost [5]. This ghost is an extra degree of freedom, apart from 5 degrees of freedom originally existing in the linear massive gravity, whose kinetic term has the wrong sign. The ghost problem had been a blockage for the massive-gravity theory until recently, in 2010, de Rham, Gabadadze, and Tolley found suitable nonlinear interactions which do not excite the Boulware–Deser mode; this is dubbed dRGT massive gravity [6,7]. Thus, massive gravity became again an active field of study.

Although it was just a generalization back then, massive gravity has its modern motivations. Introducing a non-zero mass to a graviton shrinks the scale at which the gravity works. In other words, the graviton mass weakens the gravitation at a large scale. As a result, it allows a cosmic acceleration and hence may be able to describe the mysterious dark energy in its language. This motivates cosmologists to study its cosmological implications. Moreover, since de Rham, Gabadadze, and Tolley found a healthy nonlinear massive gravity model, the theory had again opened a door to various researches on massive gravity; not only its cosmology but also the study of astrophysical objects in the theory, like black holes [8–13]. For cosmological models of massive gravity, it has been found that the solutions in the models with Minkowski fiducial metric do not admit the flat and closed FLRW solutions for the physical metric [14,15]. In order to obtain all kinds of FLRW solutions, one may consider a general form of the fiducial metric [16–20].

It has been found, however, that there are some inconsistencies when cosmology is taken into account. For example, some degrees of freedom cease to exist when the

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Friedmann–Lemaître–Robertson–Walker (FLRW) ansatz is assumed [19]. This leads to numerous studies beyond the dRGT massive gravity [21–39]. One of those is to generalize a constant graviton mass to be varied by other scalar field, dubbed mass-varying massive gravity [24–27]. The theory is proven to be free from a Boulware–Deser ghost. However, cosmological implications of such a model indicates a universe with subdominant contributions from massive gravity. In particular, the graviton mass is governed by the inverse of a scale factor of the universe which will vanish at late time. Consequently, such a model cannot give a proper explanation of the cosmic expansion caused by the massive graviton.

In this work, we propose an alternative way to construct a mass-varying massive gravity. The graviton mass is not only determined by a scalar field, but also by the kinetic term of the scalar field. Moreover, the scalar field is governed by a k-essence Lagrangian [40–42]. Under the FLRW ansatz, we found a solution whose the graviton mass do not necessarily vanish at late time. Moreover, by assuming both the k-essence and the graviton mass to behave as perfect fluids, we found that the graviton mass can give rise to a “dust-like” matter while combined with other contributions it is possible to have an equation of state parameter close to -1 , as suggested by recent observations [43]. Such matter may be responsible for a dark matter, another mysterious content known to exist in addition to the ordinary matter. Since the graviton mass can give rise to both of the dark contents, it is tempting to consider as regards its evolution whether there exists an epoch in which the two contents in the dark sector are comparable, the so-called cosmic coincidence problem.

Our paper is organized as follows. In Sect. 2, the proposed model is addressed along with its equations of motion in the FLRW background. We also discuss some crucial properties of the model in this section where we have shown the existence of the dust-like matter expected to be responsible for the dark matter. With the help of appropriate assumptions, we show in Sect. 3 the solution to this model which corresponds to the dark energy and the non-vanishing characteristic of the graviton mass existing in this model. After sketching some perspectives, we begin the dynamical system analysis in Sect. 4 to find all possible fixed points and their stabilities, and the extended analyses are covered in Sect. 5. We conclude our work in the last section by the discussion of key ideas of our work and of whether or not the coincidence problem is alleviated.

2 The model and the background equations

We consider a mass-varying dRGT massive-gravity action where the graviton mass is varied by the k-essence field. Usually, one may consider the graviton mass as a function which varies as the scalar field propagates [24–27]. However,

in this work, we will consider the graviton mass not only as a function of the scalar field ϕ but also its kinetic term $X \equiv -\frac{1}{2}g^{\mu\nu}\nabla_\mu\phi\nabla_\nu\phi$. The action of such a model can be expressed as

$$S = \int d^4x \sqrt{-g} \left[\frac{M_p^2}{2} R[g] + V(X, \phi) (\mathcal{L}_2[g, f] + \alpha_3 \mathcal{L}_3[g, f] + \alpha_4 \mathcal{L}_4[g, f]) + P(X, \phi) \right], \quad (1)$$

where R is a Ricci scalar corresponding to a physical metric $g_{\mu\nu}$, $V(X, \phi)$ is a square of the graviton mass which depends on the scalar field and its kinetic term, \mathcal{L}_i represents the interactions of the i th order of the massive graviton, and $P(X, \phi)$ is the Lagrangian of the k-essence field. In particular, those interactions of the massive graviton are constructed from two kinds of metrics and can be expressed as follows:

$$\mathcal{L}_2[g, f] = \frac{1}{2} \left([\mathcal{K}]^2 - [\mathcal{K}^2] \right), \quad (2)$$

$$\mathcal{L}_3[g, f] = \frac{1}{3!} \left([\mathcal{K}]^3 - 3[\mathcal{K}][\mathcal{K}^2] + 2[\mathcal{K}^3] \right), \quad (3)$$

$$\mathcal{L}_4[g, f] = \frac{1}{4!} \left([\mathcal{K}]^4 - 6[\mathcal{K}]^2[\mathcal{K}^2] + 3[\mathcal{K}^2]^2 + 8[\mathcal{K}][\mathcal{K}^3] - 6[\mathcal{K}^4] \right), \quad (4)$$

where the tensor $\mathcal{K}_{\mu\nu}$ is constructed from the physical metric $g_{\mu\nu}$ and an another metric $f_{\mu\nu}$ as

$$\mathcal{K}_\nu^\mu = \delta_\nu^\mu - \left(\sqrt{g^{-1}f} \right)_\nu^\mu, \quad (5)$$

where the square roots of those tensors are defined so that $\sqrt{g^{-1}f}^\mu_\rho \sqrt{g^{-1}f}^\rho_\nu = (g^{-1}f)^\mu_\nu$. In massive gravity, apart from the physical metric, there exists another kind of the metric tensor, $f_{\mu\nu}$, usually named “fiducial metric”, which is an object introduced to the theory so that one can construct non-trivial interactions from the metric tensors as in Eqs. (2), (3), and (4). Those complicated combinations in the interactions, with arbitrary values of the parameters α_3, α_4 , are to ensure the absence of the Boulware–Deser (ghostly) degree of freedom [6, 7]. Moreover, thanks to the Stuckelberg tricks, the general covariance, or the gauge symmetry, can be well integrated into the massive gravity via

$$f_{\mu\nu} = \partial_\mu \varphi^\rho \partial_\nu \varphi^\sigma \tilde{f}_{\rho\sigma}, \quad (6)$$

provided that each of the fields φ^μ transforms as a scalar under any coordinate transformation. As for the \tilde{f}_{ab} , one can choose it to be any kind of metric which shares the symmetries of the physical metric. For example, one can have a four-dimensional Minkowski metric being the fiducial metric for a cosmological solution [15], or even a higher-dimensional

kind of metric whose reduced four-dimensional metric is isotropic and homogeneous and is considered as the fiducial metric in the cosmological solution [20].

In this work, we consider the cosmological implications of the proposed model, where the isotropic and homogeneous universe is assumed whose spacetime is represented quite well by the Friedmann–Lemaître–Robertson–Walker (FLRW) metric as follows:

$$ds^2 = -N(t)^2 dt^2 + a(t)^2 \Omega_{ij}(x) dx^i dx^j, \tag{7}$$

where $N(t)$ is a lapse function, $a(t)$ represents a scale factor, which determines the scale of the spatial distance, and

$$\Omega_{ij}(\varphi) = \delta_{ij} + \frac{k \delta_{ia} \delta_{jb} \varphi^a \varphi^b}{1 - k \delta_{lm} \varphi^l \varphi^m}, \tag{8}$$

is the spatial maximally symmetric metric whose spatial curvature is characterized by $k \in \{-1, 0, +1\}$ corresponding to the open, flat, and closed geometry, respectively. As claimed, the FLRW ansatz is also used as the fiducial metric,

$$\tilde{f}_{\mu\nu} d\varphi^\mu d\varphi^\nu = -n(\varphi^0)^2 (d\varphi^0)^2 + \alpha(\varphi^0)^2 \Omega_{ij}(\varphi) d\varphi^i d\varphi^j, \tag{9}$$

where n and α are a lapse function and a scale factor in the fiducial sector. Plugging those in Eq. (1), the mini-superspace action of the model reads

$$S = \int d^4x \sqrt{\frac{1}{1 - kr^2}} \left[M_p^2 \left(-3 \frac{a\dot{a}^2}{N} + 3kNa \right) + 3Na^3 V \left(F - G \frac{n}{N} \right) + Na^3 P \right], \tag{10}$$

where

$$F \equiv \left(2 + \frac{4}{3}\alpha_3 + \frac{1}{3}\alpha_4 \right) - (3 + 3\alpha_3 + \alpha_4) \bar{X} + (1 + 2\alpha_3 + \alpha_4) \bar{X}^2 - (\alpha_3 + \alpha_4) \frac{\bar{X}^3}{3}, \tag{11}$$

$$G \equiv \frac{1}{3} (3 + 3\alpha_3 + \alpha_4) - (1 + 2\alpha_3 + \alpha_4) \bar{X} + (\alpha_3 + \alpha_4) \bar{X}^2 - \alpha_4 \frac{\bar{X}^3}{3}, \tag{12}$$

and we have defined

$$\bar{X} \equiv \frac{\alpha}{a}, \quad \eta \equiv \frac{n}{N}. \tag{13}$$

To determine the dynamics of the system, one can vary the action in Eq. (10) with respect to dynamical variables which are N, a, ϕ , and the Stuckelberg fields φ^μ . The corresponding

equations of motion, assuming the unitary gauge $\varphi^\mu = x^\mu$ for simplicity, read

$$M_p^2 \left(3H^2 + 3 \frac{k}{a^2} \right) = -3VF + 6XV_{,X} (F - G\eta) + (2XP_{,X} - P), \tag{14}$$

$$M_p^2 \left(\frac{2\dot{H}}{N} + 3H^2 + \frac{k}{a^2} \right) = -3VF + VF_{,\bar{X}} (\bar{X} - \eta) - P, \tag{15}$$

$$\frac{\dot{V}}{V} = NH (1 - h\bar{X}) \frac{F_{,\bar{X}}}{G}, \tag{16}$$

$$Na^3 (3V_{,\phi} (F - G\eta) + P_{,\phi}) = \frac{d}{dt} \left[(a^3 \sqrt{2\bar{X}}) (3V_{,X} (F - G\eta) + P_{,X}) \right], \tag{17}$$

$$3HN (-2XP_{,X} - 6XV_{,X} (F - G\eta) + VF_{,\bar{X}} (\bar{X} - \eta)) = \frac{d}{dt} (-3VF + (2XP_{,X} + 6XV_{,X} (F - G\eta)) - P), \tag{18}$$

where the last equation is obtained from the conservation on the energy-momentum tensor; $\nabla_\mu T^\mu_\nu = 0$ and we have defined

$$h \equiv \frac{H_\alpha}{H}, \quad H_\alpha \equiv \frac{\dot{\alpha}}{\alpha n}. \tag{19}$$

From the above equations, one can see that Eq. (14) is a Friedmann equation with extra matter contents coming from the graviton mass. As a partner to the Friedmann equation, the so-called acceleration equation corresponds to Eq. (15). Since we have the Bianchi identity relating the equations of motion, these five equations of motion are not entirely independent. Note that this set of equations recovers the original self-accelerating cosmology when the square of a graviton mass V is constant by which the usual condition $F_{,\bar{X}} (1 - h\eta)$ is obtained readily from Eq. (16) [15]. However, as V is no longer constant, the equations of motion look even more complex than those in general relativity. To simplify the following calculations, we choose P such that the k-essence field behaves as a perfect fluid. The appropriate form of P , which satisfies such a behavior, is

$$P(X, \phi) = P_0 X^{\frac{1+w}{2w}} = P_0 X^{\gamma/2}, \tag{20}$$

where $\gamma \equiv 2XP_{,X}/P \equiv \frac{1+w}{2w}$, P_0 is a constant, and w is an equation of state parameter corresponding to the perfect fluid represented by the k-essence field [44]. Moreover, we let the graviton mass function mimic the perfect-fluid form as

$$V = V_0 X^{\lambda/2}, \tag{21}$$

whose λ characterizes the power of the kinetic term as γ does for P , i.e. $\lambda \equiv 2XV_{,X}/V$ and V_0 is a constant. Note that under these assumptions, both P and V vary according to the kinetic term of ϕ but not the ϕ itself. Usually, in the quintessence model the continuity equation for the scalar field is obtained from the equation of motion of ϕ [45,46]. Taking that into account, we consider the equation of motion of ϕ in Eq. (17); then under the perfect-fluid assumptions for P and V in Eqs. (20) and (21) we have

$$\frac{d}{dt} \left(\left(\frac{a^3}{\sqrt{2X}} \right) (6XV_{,X} (F - G\eta) + 2XP_{,X}) \right) = 0. \quad (22)$$

After simple manipulations, the above equation gives the continuity equation for the k-essence field as

$$\frac{d}{dt} \rho_X + 3HN\rho_X = \frac{\dot{X}}{2X} \rho_X, \quad (23)$$

where we have defined

$$\rho_X \equiv (2XP_{,X} + 6XV_{,X} (F - G\eta)). \quad (24)$$

Equation (23) determines the dynamics of the matter of energy density ρ_X which resides in the Friedmann equation in Eq. (14). Interestingly, this looks exactly like a continuity equation of a “dust-like” matter with the interaction with the other matter sector determined by the flow rate of the form $\frac{\dot{X}}{2X} \rho_X$. One can also integrate Eq. (22) to find an expression for ρ_X in terms of the scale factor as

$$\rho_X = \frac{\sqrt{2X}C}{a^3}, \quad (25)$$

where C is an integration constant. In the case of a constant X , this ensures one of the properties that this matter shares with the dust; the energy density is inversely proportional to a^3 as the dust is. According to such characteristics, it is reasonable to interpret ρ_X as a dark matter. By doing so, this kind of dark matter possesses some interesting features. First of all, ρ_X is a dust-like matter which can arise naturally from the massive-gravity sector indicating that dark matter may be just an artifact of the varying graviton mass caused by the kinetic term of the k-essence field. Moreover, this claim is still valid even in the case of $P = 0$. Since a graviton mass can represent dark energy in a generic class of the dRGT massive gravity, this suggests a unification of the dark sector, namely dark energy and dark matter, by such a varying graviton mass. Second, by having this kind of matter in the theory, we may expect this model of mass-varying massive gravity to solve the cosmic coincidence problem, where the universe is known to be composed mainly of comparable amounts of dark energy and dark matter. Thanks to the unification suggested above, it may be possible to provide an

explanation on the coincidence problem by the existence of the graviton mass alone, while the cosmic acceleration also counts.

Since the equations of motion are coupled in a very cumbersome way, to get a picture of the whole of this system we need to perform a dynamical analysis, which is the main subject in the very last section. However, we can still get some rough descriptions, as a guideline to the dynamical analysis, by introducing some simple assumptions to the system, which is done in the next section.

3 Dark energy solution for the self-accelerating universe

It is widely known that our universe is expanding with an acceleration for which dark energy is responsible. There is recent observational evidence indicating that the observed effective equation of state parameter of the dark energy is close to -1 [43]. In this section, we shall adopt this characteristic by treating all the contributions from the graviton mass to have such a property. We define

$$\rho_g \equiv -3VF + 6XV_{,X} (F - G\eta), \quad (26)$$

$$p_g \equiv 3VF - VF_{,\bar{X}} (\bar{X} - \eta). \quad (27)$$

From the above definition, the corresponding equation of state parameter is defined as

$$w_g \equiv \frac{p_g}{\rho_g}. \quad (28)$$

By treating ρ_g as an energy density of dark energy, we set $w_g = -1$ and then we have the following condition:

$$6XV_{,X} (F - G\eta) = VF_{,\bar{X}} (\bar{X} - \eta). \quad (29)$$

To simplify the calculation, we use the perfect-fluid form of V in Eq. (21). Consequently, Eq. (29) becomes

$$3\lambda (F - G\eta) = F_{,\bar{X}} (\bar{X} - \eta), \quad (30)$$

$$\lambda = \frac{F_{,\bar{X}} (\bar{X} - \eta)}{3(F - G\eta)}. \quad (31)$$

Equation (31) is a requirement for the exponent λ to have a solution with the equation of state equal to -1 . To get a picture of this characteristic, let us assume

$$\bar{X} = \text{constant}, \quad (32)$$

$$\eta = \text{constant}, \quad (33)$$

$$\text{then } h = \frac{1}{\eta}. \quad (34)$$

Under these assumptions, the exponent λ in Eq. (31) is just a constant. To investigate this further, we consider Eq. (16) under the previous assumptions,

$$\begin{aligned} \frac{\dot{V}}{V} &= NH(1 - h\bar{X}) \frac{F_{,\bar{X}}}{G}, \\ \frac{\lambda \dot{X}}{2X} &= NH \left(1 - \frac{\bar{X}}{\eta}\right) \frac{F_{,\bar{X}}}{G}, \\ &= -(\bar{X} - \eta) \frac{F_{,\bar{X}}}{G\eta} \frac{\dot{a}}{a}. \end{aligned} \tag{35}$$

From the condition of λ in Eq. (31),

$$\frac{\dot{X}}{X} = -\frac{6(F - G\eta)}{G\eta} \frac{\dot{a}}{a}. \tag{36}$$

Since \bar{X} , η , and hence F and G , are constant, this equation can be integrated easily,

$$\begin{aligned} \int \frac{dX}{X} &= -\frac{6(F - G\eta)}{G\eta} \int \frac{da}{a}, \\ X &= C_0 a^{-\frac{6(F-G\eta)}{G\eta}} \end{aligned} \tag{37}$$

where C_0 is an integration constant. Now we have

$$V = V_0 X^{-\frac{(1-\frac{\bar{X}}{\eta})\eta F_{,\bar{X}}}{6(F-G\eta)}} = V_0 C_0 a \left(1 - \frac{\bar{X}}{\eta}\right) \frac{F_{,\bar{X}}}{G}. \tag{38}$$

Furthermore, Eq. (37) possibly determines a relation between the scale factor and the rate of change of the scalar field, since

$$X = \frac{\dot{\phi}^2}{2N^2} = C_0 a^{-\frac{6(F-G\eta)}{G\eta}}. \tag{39}$$

The expression of V in Eq. (38) shows the evolution of the (square of the) graviton mass as a evolves. In the previous model of mass-varying massive gravity [24–27], in which the Minkowski fiducial metric is used, the varying graviton mass shrinks as the scale factor grows. In this model, however, the exponent in Eq. (38) determines whether the graviton mass will shrink or not as the scale factor grows, or whether it will remain constant in the case that the exponent vanishes. Note that this crucial difference is caused by the different form of the fiducial metric, which is the FLRW metric in this case, to be compared with the Minkowski one in the previous models. This result indicates the sensitivity of the fiducial metric existing in the generic dRGT massive gravity where different fiducial metrics set different stages for the system and provide different solutions [16–20].

One more crucial point of this analysis is that the contributions from the graviton mass can have the same equation of state parameter as dark energy, while one of those contributions possesses the characteristic of dust, namely the

term $6XV_{,X}(F - G\eta)$. From Eq. (23), such a term belongs to the dark matter ρ_X . This may be a way out for the cosmic coincidence problem, since we may infer that varying graviton mass is responsible for a dark matter via the term like $6XV_{,X}(F - G\eta)$, as we have claimed in the previous section, while it can still drive the accelerating expansion. To verify this idea, and to seek a finer description of this model, we will perform a dynamical analysis, which can be found in the next section.

4 Dynamical system

In this section, we will consider the dynamics of the universe to be governed by this new class of mass-varying massive gravity models using the method of the autonomous system. Due to the complexity of the graviton mass, we will begin this section with a simple analysis by considering the flat FLRW where $k = 0$ and assuming that \bar{X} , η are constant over time, thus $h = 1/\eta$. From this assumption, the evolution of X is simply determined by Eq. (16) such that

$$X' = \frac{\dot{X}}{HNX} = \frac{2}{\lambda} \frac{F_{,\bar{X}}}{G} (1 - h\bar{X}) = -\frac{6s}{\lambda r}, \tag{40}$$

$$\lambda \equiv \frac{2XV_{,X}}{V}, \tag{41}$$

where the prime denotes the derivative with respect to $\ln a$. The parameters r and s are constant and defined as

$$r \equiv \frac{G\eta}{F}, \quad s \equiv \frac{F_{,\bar{X}}(\bar{X} - \eta)}{3F}. \tag{42}$$

In order to obtain a suitable autonomous system, let us define dimensionless variables as follows:

$$x = -\frac{FV}{M_p^2 H^2}, \tag{43}$$

$$z = -\frac{P}{3M_p^2 H^2}, \tag{44}$$

$$y = \frac{2XP_{,X} + 6XV_{,X}F(1 - r)}{3M_p^2 H^2} = \frac{\rho_X}{3M_p^2 H^2}, \tag{45}$$

$$\gamma \equiv \frac{2XP_{,X}}{P}. \tag{46}$$

By using these variables, the equations of motion can be written in the form of autonomous equations as

$$x' = 3x \left(y + sx - \frac{s}{r}\right), \tag{47}$$

$$y' = 3y \left(y + sx - 1 - \frac{s}{\lambda r}\right), \tag{48}$$

$$\lambda' = \frac{6s}{r} \left(\frac{\lambda}{2} - (1 + \Gamma)\right), \tag{49}$$

Table 1 Summary of the properties of the fixed points

Name	x	y	z	w_{eff}	Existence	Stability
(a)	0	0	1	-1	$\gamma = 0$	$0 \leq \frac{s}{r} \leq 1$
(b)	$\frac{1}{r}$	0	$1 - \frac{1}{r}$	$-1 + \frac{s}{r}$	$\gamma = \lambda$	$\frac{\lambda}{1-\lambda} \leq \frac{s}{r} < 0$
(c)	0	$1 + \frac{s}{\lambda r}$	$-\frac{s}{\lambda r}$	$\frac{s}{\lambda r}$	$\gamma = 1 + \frac{\lambda r}{s}$	$\frac{\lambda}{1-\lambda} < \frac{s}{r} < -1$
(d)	$\frac{1}{1+\lambda(r-1)}$	$\frac{\lambda(r-1)}{1+\lambda(r-1)}$	0	$\frac{1}{\lambda-1}$	$\lambda = \frac{s}{s-r}$	$0 < \lambda < 1$
(e)	$\frac{1+(\lambda-1)z_0}{1+\lambda(r-1)}$	$-\frac{\lambda(1-r(z_0+1))}{1+\lambda(r-1)}$	z_0	$\frac{1}{\lambda-1}$	$\lambda = \gamma = \frac{s}{s-r}$	$0 < \lambda < 1$

$$1 = x + y + z, \tag{50}$$

$$y = -\lambda x(1 - r) - z\gamma, \tag{51}$$

where $\Gamma \equiv XV_{,XX}/V_{,X}$. Since we have five variables with two constraints, it is sufficient to consider only three equations. Note that the constraint in Eq. (50) is derived from Eq. (14), while the constraint in Eq. (51) is obtained from the definition of y in Eq. (45). The equation of λ in Eq. (49) is not directly dependent on the other variables. Therefore, in principle, we can solve it separately. For simplicity, we can consider λ as a parameter and then consider only the autonomous equations with two variables, x and y . We will extend our analysis to a more general case with λ being the variable in the next section. The effective equation of state parameter can be written in terms of the dimensionless variables as

$$w_{\text{eff}} = \frac{P + 3VF - VF_{,\bar{X}}(\bar{X} - \eta)}{3M_p^2 H^2} = -z - x + xs = -1 + y + xs. \tag{52}$$

From these autonomous equations, the corresponding fixed points can be found by evaluating $x' = 0$ and $y' = 0$ in Eqs. (47) and (48), respectively. The properties of all the fixed points are summarized in Table 1, while the analyses are separately discussed for each of the fixed points below.

4.1 Fixed point (a)

From Eqs. (47) and (48), it is obvious that the system has a fixed point $(x, y) = (0, 0)$. By using the constraint equations, one obtains $z = 1$ and $\gamma = 0$. This means that the function P is constant and then this point corresponds to general relativity with a cosmological constant where the universe is dominated by the cosmological constant. To ensure such a claim, one can compute the corresponding effective equation of state parameter, which yields $w_{\text{eff}} = -1$. This is exactly the equation of state parameter of the cosmological constant which drives the accelerating de Sitter expansion.

The stability of the fixed point can be found by analyzing the eigenvalues of the linearly perturbed autonomous equations. By performing the linear perturbations, the eigenvalues can be written as $(\mu_1, \mu_2) = (-3s/r, -3 - 3s/r)$. The

stability requires both of the eigenvalues to be negative, or otherwise the fixed point is said to be unstable or to be a saddle fixed point. In this case, the signs of those eigenvalues are determined by the value of the term $\frac{s}{r} = (\bar{X} - \eta) \frac{F_{,\bar{X}}}{G\eta}$, which means $0 \leq \frac{s}{r} \leq 1$ for the stable fixed point. Note that in the case of vanishing eigenvalues, like $s = 0$, one has to consider the perturbations up to second order or use a numerical investigation in order to determine the stability. In this analysis, we ensure the stability in this case by the numerical method and we have found that it is stable.

Even though this fixed point can provide a period of late-time expansion, it is not much of interest due to the disappearance of the graviton mass. This resulting property is one of the drawbacks in the previous model of mass-varying massive gravity [24–27].

4.2 Fixed point (b)

One of possible fixed points may be in the form $(x, y) = (x_0, 0)$ by which the universe is governed mainly by massive gravity alone. From Eq. (47), one can find x_0 as follows:

$$x_0 = \frac{1}{r}. \tag{53}$$

According to Eq. (45), there are two possible solutions for this kind of fixed point. One is $r = 1$ in which $x_0 = 1, z_0 = 0$, and another one is $\lambda = \gamma$ in which $x_0 = \frac{1}{r}, z_0 = 1 - \frac{1}{r}$. The effective equation of state parameter can be written as

$$w_{\text{eff}} = -1 + \frac{F_{,\bar{X}}(\bar{X} - \eta)}{3G\eta} = -1 + \frac{s}{r}. \tag{54}$$

Interestingly, $w_{\text{eff}} = -1$ as $F_{,\bar{X}} = 0$ or $(\bar{X} - \eta) = 0$. This characteristic is a usual cosmological solution of the original massive gravity. In particular, this condition indicates that the graviton mass ceases to vary, according to Eq. (16). Moreover, since in this case $z = 1 - \frac{1}{r}$, the pressure of the k-essence field is non-zero for $r > 1$, which means the k-essence field is supposed to be a form of matter with non-zero pressure (not dust).

In order to find the stability condition for this fixed point, one can find the eigenvalues of the linearly perturbed autonomous equations, which can be written as

$$(\mu_1, \mu_2) = \left(3\frac{s}{r}, -3 + 3\frac{(\lambda - 1)s}{\lambda r} \right). \tag{55}$$

Again, both of the eigenvalues contain the term s/r , and then the fixed point will be stable if $\frac{\lambda}{1-\lambda} \leq \frac{s}{r} < 0$. Note that, for this fixed point, it is possible to provide $w_{\text{eff}} < -1$ to satisfy the observation, which indicates that the mean value of the equation of state parameter is slightly less than -1 [43].

4.3 Fixed point (c)

One can obtain a fixed point such that $(x, y) = (0, y_0)$. From Eq. (47), one can find y_0 as follows:

$$y_0 = 1 + \frac{s}{\lambda r}. \tag{56}$$

By using the constraint equation in Eq. (50), one obtains $z_0 = -\frac{s}{\lambda r}$. From the constraint equation in Eq. (51), we have

$$\gamma = -\frac{y}{z} = -1 + \frac{1}{z} = 1 + \frac{1}{w_m}, \tag{57}$$

where w_m is the equation of state parameter of the fluid contributed from $P(X) = P_0 X^{(1+w_m)/2w_m}$. The effective equation of state parameter can be written as

$$w_{\text{eff}} = -z = \frac{s}{\lambda r}. \tag{58}$$

Again, there exist two significant branches of the solution such that this fixed point is a matter-dominated point. If $z = 0$, this corresponds to $w_{\text{eff}} = 0$, which leads to the universe being in a matter-dominated period.

The eigenvalues of the autonomous system can be written as

$$(\mu_1, \mu_2) = \left(3 + 3\frac{s}{\lambda r}, 3 - 3\frac{s(\lambda - 1)}{\lambda r} \right). \tag{59}$$

If one requires this point to represent the matter-dominated epoch, one must set the parameters so that this point is unstable. This means the universe should evolve through this point to end up in other stable points since we know the matter-dominated epoch should exist in the universe’s timeline but not nowadays. One can see that, for small negative value of s/r , the universe can evolve in the standard history at which fixed point (c) corresponds to a matter-dominated period with $w_{\text{eff}} \sim 0$, and fixed point (b) corresponds to the late-time expansion of the universe due to the contribution from the graviton mass. However, it is not possible to alleviate the coincidence problem, since the contribution of non-relativistic matter vanishes at late time.

4.4 Fixed point (d)

According to Eqs. (47) and (48), one may consider the fixed point corresponding to the non-zero x and y . This point can

be obtained by evaluating both (non-zero) x and y from Eqs. (50), (51), and (47), while a constraint on the parameters by which the non-zero (x, y) exist can be obtained from Eqs. (47) and (48). After simple manipulation, we have

$$x = \frac{1}{1 + \lambda(r - 1)}, \quad y = \frac{\lambda(r - 1)}{1 + \lambda(r - 1)}, \quad \text{and } z = 0, \tag{60}$$

where γ is arbitrary and λ is fixed to be $\lambda = \frac{s}{s-r}$. The effective equation of state parameter can be written as

$$w_{\text{eff}} = \frac{1}{\lambda - 1}. \tag{61}$$

To determine the stability of this point, we find the eigenvalues of the system of equations. Interestingly, this point renders the two autonomous equations degenerate. This can be seen by computing the linear perturbed equations for both x and y evaluated at this fixed point. The eigenvalues of this autonomous system are expressed as

$$(\mu_1, \mu_2) = \left(0, \frac{3\lambda}{\lambda - 1} \right). \tag{62}$$

The vanishing eigenvalue here is nothing but an artifact of the degeneracy due to this fixed point. In particular, it is possible to redefine the variables such that the problem is reduced into a one-dimensional system. With such a redefinition, the stability of this fixed point is due to the non-zero eigenvalue in Eq. (62), which can be negative when $0 < \lambda < 1$. If this condition is taken into account, requiring the fixed point (c) to represent the matter-dominated era will restrict the combination $\frac{s}{r}$ to vanish.

This fixed point seems to provide a possible way to alleviate the coincidence problem due to the non-zero y . However, it cannot be used since, at the late-time expansion, w_{eff} must approach -1 and then lead to the fact that $(x, y) \rightarrow (1, 0)$. Nevertheless, it still provides an interesting result. For the case of $s = 0$ and $0 < \lambda \ll 1$, this fixed point is stable, while the fixed point (b) is unstable and then we can use this fixed point as the one for the late-time expansion of the universe. For this condition the fixed point (c) is still used for the matter-dominated period with $z = 0$. Therefore, this means that it is possible to obtain $z = 0$ for the whole history of the universe. This leads to the fact that, without providing an extra non-relativistic matter field such as dark matter, the contribution from the graviton mass can play the role of both dark matter and dark energy. This is one of the crucial properties of this model, since it can unify the two main unknown contents of the universe, dark matter and dark energy, by using only a graviton mass.

4.5 Fixed point (e)

Similarly to the derivation in fixed point (d), one can solve an algebraic equation by imposing $\gamma = \lambda$ and requiring non-zero x and y . As the result, the fixed point can be expressed as

$$x = \frac{1 + (\lambda - 1)z_0}{1 + \lambda(r - 1)}, \quad y = -\frac{\lambda(1 - r(z_0 + 1))}{1 + \lambda(r - 1)}, \quad z = z_0, \tag{63}$$

where $\gamma = \lambda = \frac{s}{s-r}$ and z_0 is arbitrary. The effective equation of state parameter is the same as the one in the fixed point (d), which can be written as

$$w_{\text{eff}} = \frac{1}{\lambda - 1}. \tag{64}$$

Moreover, the eigenvalues for the stability analysis are still the same as for the fixed point (d) and then the stability condition for this fixed point can be expressed as $0 < \lambda < 1$. Even though this fixed point shares most properties with fixed point (d), it cannot provide the unification of the two dark components, since z must have a non-zero value.

From the above analyses, we experienced the incompatibility between matter domination and the present dark energy domination. One may see that for a large λ , the fixed point (c) can represent the matter-dominated epoch, while the small value of λ is needed in the fixed point (d) or (e) to solve the coincidence problem. It is natural to generalize the theory further by allowing λ to change appropriately in time. This idea will be adopted and carefully analyzed in the next section.

5 Extended analyses

As we have mentioned, even though the model can be used to unify the dark contents of the universe, it still cannot be used to solve the coincidence problem. According to our analysis, this is due to the fact that λ is set to be a constant. In this section, we will show the possibility to solve the coincidence problem when λ is set as a dynamical variable. For completeness, we will add radiation into our consideration and then use numerical method to show that the radiation does not affect the unification in the dark sector. Note that the equation of motion for the radiation is obtained by using the conservation of its energy-momentum tensor or the continuity equation. By including the radiation and taking λ as a dynamical variable, the autonomous equations can be written as

$$x' = 3x \left(y + sx - \frac{s}{r} + \frac{4}{3}\Omega_r \right), \tag{65}$$

$$y' = 3y \left(y + sx - 1 - \frac{s}{\lambda r} + \frac{4}{3}\Omega_r \right), \tag{66}$$

$$\Omega_r' = 3\Omega_r \left(y + sx + \frac{4}{3}(\Omega_r - 1) \right), \tag{67}$$

$$\lambda' = \frac{6s}{r} \left(\frac{\lambda}{2} - (1 + \Gamma) \right), \tag{68}$$

$$1 = x + y + z + \Omega_r, \tag{69}$$

$$y = -\lambda x(1 - r) - zy, \tag{70}$$

$$\Omega_r \equiv \frac{\rho_r}{3M_p^2 H^2}, \tag{71}$$

where ρ_r is the energy density of the radiation. The effective equation of state parameter can be written as

$$w_{\text{eff}} = -1 + y + xs + \frac{4}{3}\Omega_r. \tag{72}$$

From Eq. (67), we can see that all fixed points we found in the previous section still exist with $\Omega_r = 0$. Also, there exists the unstable fixed point such that $\Omega_r = 1$, while x and z (hence y) vanish. From Eq. (68), one can see that λ does not couple to the others and the fixed point takes place at $\lambda = 2(\Gamma + 1)$. For simplicity, one can set Γ as a constant. In order to confirm the claim in the previous section that there exists a standard evolution without introducing a k-essence Lagrangian or in the case of $z = 0$, we use numerical methods to evaluate the equations above by setting $s = 0$. The evolutions of x , y , and Ω_r are illustrated in the left panel of Fig. 1, and the evolution of the effective equation of state parameter is shown in the right panel of Fig. 1. We can see that there exists non-relativistic matter, inferred as dark matter represented by the variable y , while the variable x represents the dark energy that drives the late-time expansion of the universe. Both x and y are contributed from the graviton mass.

Now, let us consider the possibility to solve the coincidence problem. Let us use the fixed point (e) to be one corresponding to the late-time expansion of the universe. For this fixed point, the parameters s , r , and Γ are obtained by giving the initial conditions for the dynamical variables. In order to obtain the dynamics of all variables, we have to put the initial conditions slightly away from the fixed point. It is sufficient to put λ slightly above the fixed point, since we need λ to grow as time goes backward to ensure that it will have a high enough value for the matter-dominated period. In order to obtain $w_{\text{eff}} \sim -1$ at the present time, we have to set the value of the variable λ at the fixed point as $\lambda_f \rightarrow 0$. As a result, $\frac{s}{r} = \frac{\lambda_f}{\lambda_f - 1} \rightarrow 0$. In order to obtain a proper matter-dominated period, one has to put the initial value of λ far away from the fixed point. This situation makes the fixed point (b) stable and then the system evolves to the point (b)

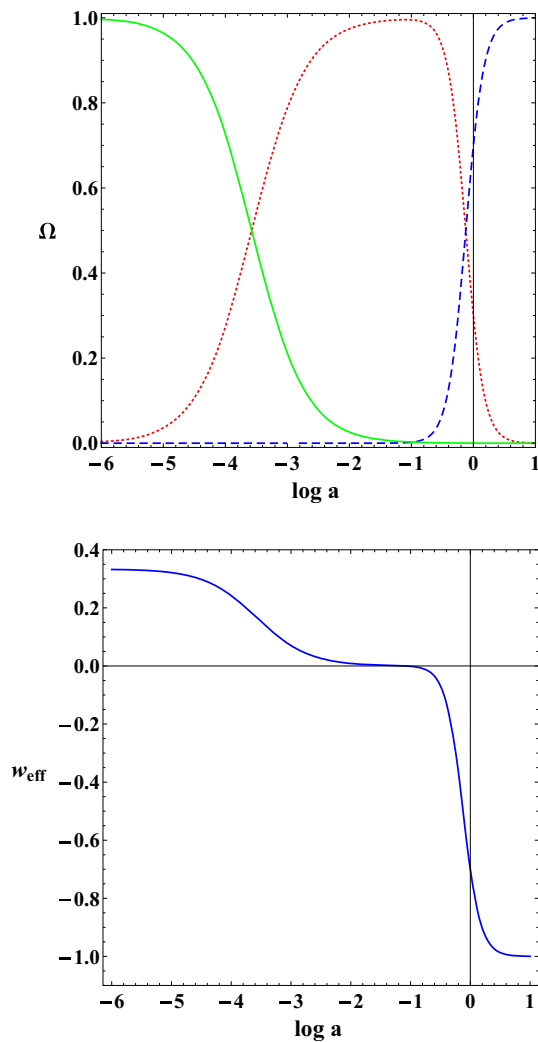


Fig. 1 The left panel shows the evolution of x , y , and Ω_r . The dotted-red line represents the evolution of x , the dashed-blue line represents the evolution of x , and the solid-green line represents the evolution of Ω_r . For the right panel, the evolution of w_{eff} is represented

eventually. Therefore, in order to have the fixed point (e) at late time, one has to set w_{eff} below -1 at the fixed point, so that the point (e) becomes a stable point. According to this setting, we show the evolution of the dynamical variables reaching the fixed point (e) to alleviate the coincidence problem in Fig. 2. Note that we set $\lambda_f = 0.4$, leading to $w_{\text{eff}} = -1.67$ and $\lambda_0 = 1.0$.

In order to overcome the incompatibility among the fixed points, one may extend the analysis by allowing s , Γ or r to be dynamical variables. This will make the dynamical system more complicated. We found another possibility to overcome this incompatibility by imposing the constraint $\lambda = \gamma$ for the entire evolution. As a result, we have only three independent equations for six variables and three constraints. The dynamical variable λ can be written in terms of other variables as

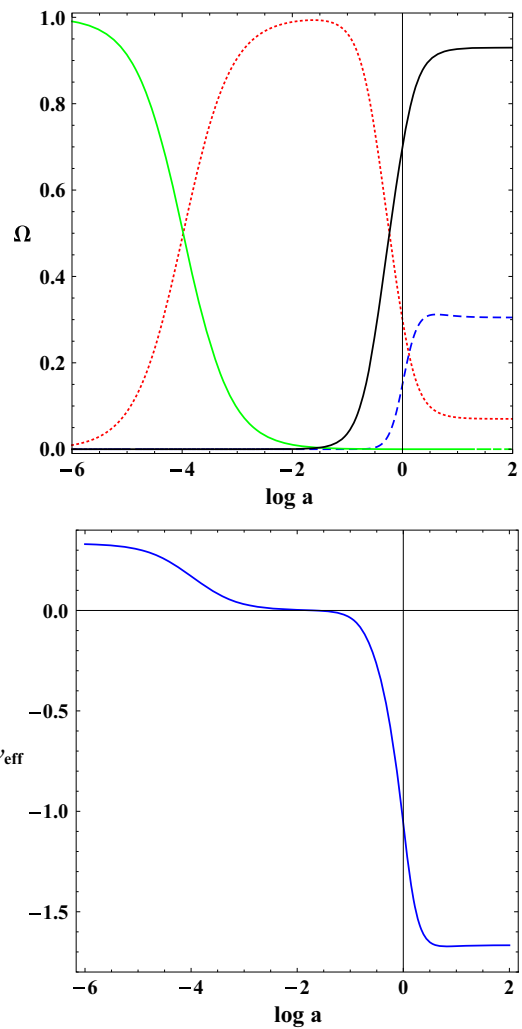


Fig. 2 The left panel shows the evolution of x , y , $x + z$, and Ω_r . The dotted-red line represents the evolution of x , the dashed-blue line represent the evolution of x , the solid-black line represents the evolution of $x + z$ and the solid-green line represents the evolution of Ω_r . For the right panel, the evolution of w_{eff} is represented. We set the parameters such that $\lambda_f = 0.4$ and $\lambda_0 = 1.0$ where λ_f is the value at the fixed point and λ_0 is one at the present time

$$\lambda = \frac{y}{rx + y + \Omega_r - 1}. \tag{73}$$

As a result, by setting the initial condition at the radiation dominated period, the evolution of the dynamical variables and the effective equation of state are shown in Fig. 3. From this figure, one can see that the evolution of the universe reaches the fixed point (e) at late time while the matter and radiation period are also properly presented. For the plot in this figure, we set $\lambda_f = 0.02$, and then the consequent results are $\Gamma = -0.99$ and $w_{\text{eff}} \sim -1.02$. Note that the behavior of the resulting plot in Fig. 3 is sensitive to the initial value of x at the radiation dominated period where we set it choosing $x_i \sim 10^{-16}$.

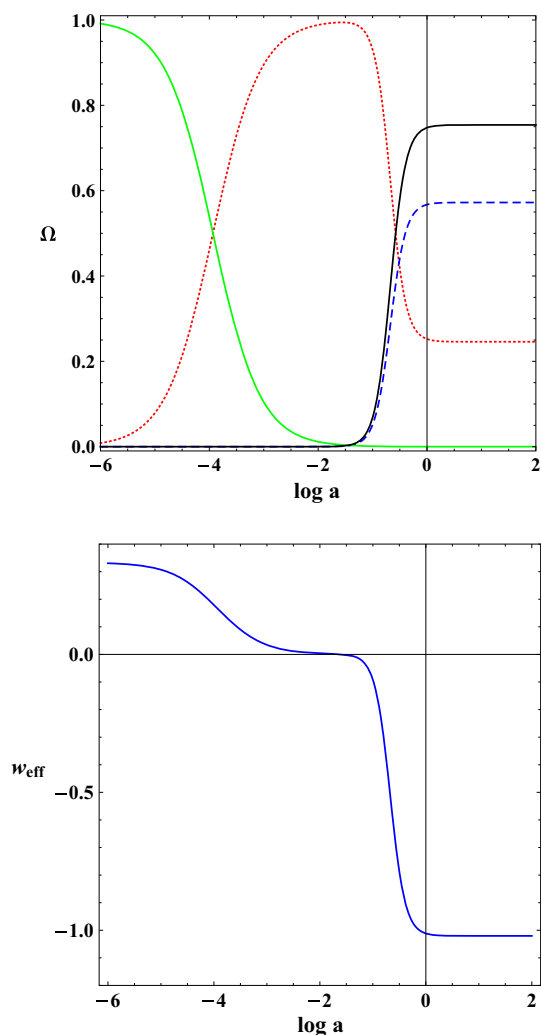


Fig. 3 The left panel shows the evolution of x , y , $x + z$, and Ω_r . The dotted-red line represents the evolution of x , the dashed-blue line represent the evolution of x , the solid-black line represents the evolution of $x + z$, and the solid-green line represents the evolution of Ω_r . For the right panel, the evolution of w_{eff} is presented

6 Conclusion

We have constructed a new class of mass-varying massive gravity models, in which not only the k-essence field but also its kinetic term determines the variation of the graviton mass. We have shown in Sect. 2 that there is a possibility for the graviton mass to live at late time compared with the previous model whose the graviton mass only depends on the scalar field and shrinks as the universe grows [24–27]. After simple manipulations and under particular assumptions, we found that a “dust-like” matter which behaves like a non-relativistic dust can naturally result from the graviton mass and it is a possible candidate for dark matter. This can be seen more clearly in the case $P = 0$ in which the dark matter comes solely from the varying graviton mass. Having such matter in the system, this model of massive gravity can describe

the cosmic accelerating expansion with the equation of state parameter close to -1 , while the universe is not entirely dominated by the dark energy part contributed also by the graviton mass. This property signals a possibility of having the universe composed of comparable amounts between dark energy and dark matter, known as the cosmic coincidence problem. To obtain a finer description on this, the usual method of the dynamical analysis is performed by taking the dark matter candidate into account and the results are carefully investigated as regards the issue of the coincidence problem. For the first simple case, the exponent of the kinetic term in the graviton mass λ is kept constant. We found the fixed points which correspond to various epochs in the history of the universe such as the matter-dominated period and massive-gravity-dominated periods. However, to have those fixed points with the appropriate stabilities in the evolution of the universe, the results suggest a system with λ as additional variable. The more general case, where λ is allowed to vary, is investigated where the radiation is included. While the result covers all the fixed points in the constant λ case, this allows the evolution in which there exists a matter-dominated period as well as a late-time expansion epoch. There are several crucial points in this investigation. First, we obtain the universe in which the graviton mass serves as both dark energy and dark matter, while it can still drive the cosmic acceleration. Second, to solve the coincidence problem, we obtain a universe with the effective equation of state parameter significantly below -1 unless both λ and γ are set equal with one another for the entire evolution of the universe. Since the analyses are under particular assumptions, this model still has room for study in more complicated ways. For example, one can exclude the assumptions proposed in this work for a more complex system or one can consider this model in a different aspect, like its astrophysical implications. Not only as regards the applications, but also studying the theoretical consistency, whether there exists a ghost instability or not, is a worthy challenge which we leave for future work. Apart from the constraints mentioned, one may think of constraining the model with various observations. This idea is also interesting, since the observations may judge the fate of this model by tightening it with constraints.

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