

The new ν MSM ($\nu\nu$ MSM): radiative neutrino masses, keV-scale dark matter and viable leptogenesis with sub-TeV new physics

Sven Baumholzer,^a Vedran Brdar^b and Pedro Schwaller^a

^a*PRISMA Cluster of Excellence and Mainz Institute for Theoretical Physics,
Johannes Gutenberg-Universität Mainz,
Staudingerweg 9, 55099 Mainz, Germany*

^b*Max-Planck-Institut für Kernphysik,
Saupfercheckweg 1, 69117 Heidelberg, Germany*

E-mail: sbaumhol@students.uni-mainz.de, vbrdar@mpi-hd.mpg.de,
pedro.schwaller@uni-mainz.de

ABSTRACT: We study the scenario in which the Standard model is augmented by three generations of right-handed neutrinos and a scalar doublet. The newly introduced fields share an odd charge under a \mathbb{Z}_2 parity symmetry. This model, commonly known as “Scotogenic”, was designed to provide a mechanism for active neutrino mass generation as well as a viable dark matter candidate. In this paper we consider a scenario in which the dark matter particle is at the keV-scale. Such particle is free from X-ray limits due to the unbroken parity symmetry that forbids the mixing between active and right-handed neutrinos. The active neutrino masses are radiatively generated from the new scalars and the two heavier right-handed states with $\sim \mathcal{O}(100)$ GeV masses. These heavy fermions can produce the observed baryon asymmetry of the Universe through the combination of Akhmedov-Rubakov-Smirnov mechanism and recently proposed scalar decays. To the best of our knowledge, this is the first time that these two mechanisms are shown to be successful in any radiative model. We identify the parameter space where the successful leptogenesis is compatible with the observed abundance of dark matter as well as the measurements from the neutrino oscillation experiments. Interestingly, combining dark matter production and successful leptogenesis gives rise to strict limits from big bang nucleosynthesis which do not allow the mass of dark matter to lie above ~ 10 keV, providing a phenomenological hint for considered low-scale dark matter. By featuring the keV-scale dark matter free from stringent X-ray limits, successful baryon asymmetry generation and non-zero active neutrino masses, the model is a direct analogue to the ν MSM model proposed by Asaka, Blanchet and Shaposhnikov. Therefore we dub the presented framework as “The new ν MSM” abbreviated as $\nu\nu$ MSM.

KEYWORDS: Beyond Standard Model, Cosmology of Theories beyond the SM

ARXIV EPRINT: [1806.06864](https://arxiv.org/abs/1806.06864)

Contents

1	Introduction	1
2	The model	3
2.1	Relevant constraints	3
2.2	Active neutrino masses	5
3	Dark matter production	6
3.1	Production via scalar decays	7
3.2	Production via N_2 decays	9
4	Leptogenesis	12
5	Identifying the viable joint parameter space for DM and leptogenesis	19
6	Detection prospects	22
7	Summary and conclusions	23
A	Phase space integration and reaction densities	24

1 Introduction

The Standard model (SM) is a remarkably accurate theory. Its particle content and interactions are almost flawlessly mapping Nature’s choice. However, there are still several open questions and in particular those that stand out are:

- What is the origin of neutrino mass?
- Does dark matter (DM) interact with SM particles?
- Through which mechanism was the observed asymmetry between matter and anti-matter generated?

The indisputable answer to any of these questions would indicate a tremendous progress for physics, in particular for the community striving to discover “new physics”, as it is by now clear that the solution does not lie within the SM.

We attempt to address all of the above questions within the model proposed by Ma [1], dubbed “Scotogenic”, in which the SM field content is supplemented by three right-handed neutrinos and a scalar doublet, all odd under a postulated \mathbb{Z}_2 parity symmetry. The model was originally envisioned to account for small active neutrino masses and the electroweak scale DM produced via standard thermal freeze-out. Instead, we put in focus keV-scale for

the mass of the fermionic DM particle. We consider two viable regimes for DM production and ensure that the studied parameter space is fully consistent with leptonic mixing parameters and the observed neutrino mass squared differences.

In addition, we scrutinize the generation of BAU (baryon asymmetry of the Universe) from the produced lepton number asymmetry (leptogenesis). Such asymmetry transfer is viable due to the existence of non-perturbative sphaleron processes at high temperatures. We consider two complementary mechanisms, namely BAU production from right-handed neutrino oscillations introduced by Akhmedov, Rubakov and Smirnov (ARS) [2] as well as recently proposed asymmetry generation from scalar decays [3].

The motivation for this work stems from the ν MSM model proposed by Asaka, Blanchet and Shaposhnikov [4–6]. In the ν MSM, the SM particle content is extended with only three right-handed neutrinos, where the lightest one is a keV-scale DM produced via neutrino oscillations [7, 8] due to the mixing between active and right-handed neutrinos. The heavier two GeV-scale right-handed states generate active neutrino masses in the seesaw type-I model [9–12] as well as produce BAU through the ARS mechanism. When confronted with the current experimental data, the ν MSM is seriously challenged. In particular, the vast portion of the viable parameter space for DM is excluded by the combination of structure formation and X-ray limits [13–15].

A keV-scale DM candidate and ARS mechanism for baryogenesis are also prominent characteristics in our scenario. Furthermore, the relevant fermionic Yukawa and mass terms in these two models differ only in the employed scalar doublet — the SM Higgs doublet is considered in ν MSM whereas our model hinges on the existence of a second \mathbb{Z}_2 -odd scalar doublet. In contrast with ν MSM, in the Scotogenic setup keV-scale DM does not mix with the active sector due to the imposed \mathbb{Z}_2 symmetry. Hence, the DM parameter space opens up due to the absence of astrophysical X-ray limits [16–19]. Let us note that, in such a framework, the controversial 3.5 keV line [20–24] does not have a DM origin and the atomic physics explanation is favored [19, 25–27].

As in the original ν MSM, we successfully identify the parameter space corresponding to non-zero neutrino masses, correct DM relic abundance and the observed amount of baryon asymmetry in the Universe. Hence, we simultaneously address all three of the aforementioned questions. The strongest constraint on our model, arising from the measurements of primordial abundances of light nuclei, forbids the DM candidate to be heavier than ~ 10 keV.

The structure of the paper is as follows. In section 2 we introduce the model and discuss relevant theoretical and experimental limits. In particular, we show how the neutrino masses are generated and discuss conditions which ensure the correct low-energy neutrino phenomenology (mixing angles and mass squared differences) in our numerical scans. In sections 3 and 4 we identify the viable parameter space for DM and BAU in the model, respectively. Results from these two sections are combined in section 5 with the purpose of finding the regions where DM and BAU can be addressed simultaneously. In section 6 we discuss the implications from structure formation and feasibility of probing this model at various experimental facilities. Finally, in section 7 we conclude.

2 The model

We supplement the Standard Model particle content with an additional scalar doublet $\Sigma = (\sigma^+, \sigma^0)^T$, and three right-handed neutrinos N_i . We study the spectrum in which all new scalars are heavier than the right-handed neutrinos, of which the lightest one, N_1 , is a keV-scale DM candidate. We assume that all these newly introduced degrees of freedom have an odd ($-$) charge under a \mathbb{Z}_2 parity symmetry, whereas the SM particles have an opposite, even ($+$) charge. The scalar sector is therefore equivalent to the one in the inert doublet model [28] with the potential that reads [1, 29]

$$V = \mu_1^2 \Phi^\dagger \Phi + \mu_2^2 \Sigma^\dagger \Sigma + \frac{1}{2} \lambda_1 (\Phi^\dagger \Phi)^2 + \frac{1}{2} \lambda_2 (\Sigma^\dagger \Sigma)^2 + \lambda_3 (\Phi^\dagger \Phi)(\Sigma^\dagger \Sigma) + \lambda_4 (\Phi^\dagger \Sigma)(\Sigma^\dagger \Phi) + \frac{\lambda_5}{2} \left((\Phi^\dagger \Sigma)^2 + \text{h.c.} \right), \quad (2.1)$$

where $\Phi = (\phi^+, \phi^0)^T$ is the SM Higgs doublet containing the Higgs boson with mass equal to $2\lambda_1 v^2$, where $v = 246/\sqrt{2}$ GeV denotes the vacuum expectation value of ϕ^0 .

By introducing an additional doublet, the scalar sector contains four additional degrees of freedom with masses

$$\begin{aligned} m_\pm^2 &= \mu_2^2 + \lambda_3 v^2, \\ m_S^2 &= \mu_2^2 + (\lambda_3 + \lambda_4 + \lambda_5) v^2, \\ m_A^2 &= \mu_2^2 + (\lambda_3 + \lambda_4 - \lambda_5) v^2. \end{aligned} \quad (2.2)$$

Here, m_\pm , m_S and m_A are the masses of the charged, CP-even and CP-odd scalar, respectively. In the remainder of the paper we will denote charged scalars with σ^\pm and neutral CP-even (CP-odd) scalar with S (A).

In the fermion sector, the presence of a \mathbb{Z}_2 symmetry forbids the “traditional” lepton portal

$$y_\phi \bar{N} \tilde{\Phi}^\dagger L + \text{h.c.}, \quad (2.3)$$

where L is the SM lepton doublet and y_ϕ is a corresponding Yukawa coupling. However, the Yukawa interaction between N_i , leptons and Σ field is allowed. After adding a Majorana mass term for right-handed neutrinos, the relevant lepton sector Lagrangian reads

$$\mathcal{L} \supset y_{i\alpha} \bar{N}_i \tilde{\Sigma}^\dagger L_\alpha + \frac{1}{2} m_{N_i} \bar{N}_i N_i^c + \text{h.c.}, \quad (2.4)$$

where α denotes the SM lepton generations and m_{N_i} is the mass of i -th right-handed neutrino. Without loss of generality we take the right-handed neutrino mass matrix in the diagonal form. Note that the replacement $\Sigma \rightarrow \Phi$ transforms the Yukawa term in eq. (2.4) into the one given in eq. (2.3).

2.1 Relevant constraints

Before discussing a realization of nonzero neutrino masses we present the most relevant theoretical and experimental constraints for the considered model.

- *Big Bang Nucleosynthesis.*

The most relevant limit arises from the measurement of primordial abundances of light nuclei [30–32], known as the big bang nucleosynthesis (BBN). These precisely measured values are sensitive to energy injection into the plasma, for example from the decays of long-lived particles during the BBN epoch. In our model, late decays of N_2 to keV-scale DM (N_1) could spoil the BBN predictions. The decay rate for such process is [33]

$$\Gamma(N_2 \rightarrow N_1 l_\alpha^\pm l_\beta^\mp) = \frac{m_{N_2}^5}{6144\pi^3 m_\pm^4} (|y_{1\alpha}|^2 |y_{2\beta}|^2 + |y_{1\beta}|^2 |y_{2\alpha}|^2), \quad (2.5)$$

where l_α are the charged SM leptons of generation α . The BBN limit is relevant chiefly due to the tiny $y_{1\alpha}$ couplings, necessary to keep DM out of the thermal equilibrium with SM bath. The findings from the recent analysis presented in ref. [32] allow us to infer the masses and abundances of N_2 consistent with respect to BBN constraint. A quantitative discussion of the impact of these constraints on our model parameter space is presented in section 5.

- *Structure Formation.*

Due to the absence of mixing between active and right-handed neutrinos, X-ray limits [14] on keV-scale DM are non-existent within the presented model which opens up a viable parameter space for light right-handed neutrinos. However, one also needs to take into consideration the parameter space excluded by structure formation bounds (*Lyman- α* forests) and Milky Way satellite counts [34]. Throughout this paper, we will show results consistent with the limits arising from the most conservative scenario where the keV-scale DM is assumed to inherit a thermal distribution function with $\langle p \rangle / T \approx 3.1$. In section 6 we discuss how this spectrum can be made colder. From refs. [15, 34–37] we infer that in order to be in accord with the Milky Way satellite count, $m_{N_1} \gtrsim 6$ keV is viable. The existing limits from *Lyman- α* forests are stronger but rather controversial due to effects stemming from inter-galactic medium [37, 38] and therefore we do not adopt them.

- *Lepton Flavor Violation and Scalar Potential.*

The scotogenic model predicts lepton flavor violation processes [39] of the type $l_\alpha \rightarrow l_\beta \gamma$ and $l_\alpha \rightarrow 3l_\beta$. In table 1 we summarize the bounds given in ref. [40].

The quartic couplings in the scalar potential receive constraints from the requirement of the vacuum stability [29, 41]

$$\lambda_3 > -\sqrt{\lambda_1 \lambda_2}, \quad \lambda_3 + \lambda_4 - |\lambda_5| > -\sqrt{\lambda_1 \lambda_2}, \quad (2.6)$$

where $\lambda_1 > 0$ and $\lambda_2 > 0$.

We find the other constraints [42], arising for instance from the compatibility with the electroweak precision data (Peskin-Takeuchi parameters) [43], to be much weaker with respect to the aforementioned ones.

LFV process	BR upper bound
$\mu^+ \rightarrow e^+ \gamma$	$4.2 \cdot 10^{-13}$
$\tau^\pm \rightarrow e^\pm \gamma$	$3.3 \cdot 10^{-8}$
$\tau^\pm \rightarrow \mu^\pm \gamma$	$4.4 \cdot 10^{-8}$
$\mu \rightarrow 3e$	$1.0 \cdot 10^{-12}$
$\tau \rightarrow 3e$	$2.7 \cdot 10^{-8}$
$\tau \rightarrow 3\mu$	$2.1 \cdot 10^{-8}$

Table 1. Constraints on LFV processes, taken from [40]. The left column indicates rare LFV processes and the right one shows the corresponding upper limits on the branching ratios (BR).

2.2 Active neutrino masses

The Scotogenic model is known as one of the most minimal realizations [44–52] of non-vanishing neutrino masses established at oscillation experiments [53, 54]. In order to obtain the general formula for the neutrino mass matrix in the flavor basis, it is required to calculate the self-energy contributions to the active neutrino propagator, arising from the exchange of both S and A . The Majorana neutrino mass matrix reads [1, 39]

$$(m_\nu)_{\alpha\beta} = \sum_i \frac{y_{i\alpha} y_{i\beta} m_{N_i}}{16\pi^2} \left[\frac{m_S^2}{m_S^2 - m_{N_i}^2} \ln \left(\frac{m_S^2}{m_{N_i}^2} \right) - \frac{m_A^2}{m_A^2 - m_{N_i}^2} \ln \left(\frac{m_A^2}{m_{N_i}^2} \right) \right] \equiv \sum_i y_{i\alpha} y_{i\beta} \Lambda_i, \quad (2.7)$$

where the summation index i denotes the right-handed neutrino generations. For later convenience, we introduced Λ which abbreviates all the mass matrix components apart from the Yukawa couplings. As we will demonstrate in section 3, in order not to overclose the Universe, the Yukawa couplings of N_1 ($y_{1\alpha}$) must be suppressed with respect to the second and third generation ones. Therefore, N_1 does not effectively yield a contribution to eq. (2.7) which consequently sets the lightest active neutrino to be massless.

For obtaining the Yukawa couplings that are in accord with the active neutrino mass squared differences and mixing parameters, we employ the Casas-Ibarra parametrization [55]. Phenomenologically, following the above discussion on the relevance of only N_2 and N_3 states for active neutrino mass generation, the corresponding 2×3 Yukawa submatrix reads

$$Y_{23} = \left(\sqrt{\Lambda_{23}} \right)^{-1} R \sqrt{m_\nu^{\text{diag}}} U_{\text{PMNS}}^\dagger. \quad (2.8)$$

Here, $\Lambda_{23} = \text{diag}(\Lambda_2, \Lambda_3)$ and R is an orthogonal matrix parametrized with one complex angle

$$R = \begin{pmatrix} 0 & \cos(\omega - i\xi) & -\sin(\omega - i\xi) \\ 0 & \sin(\omega - i\xi) & \cos(\omega - i\xi) \end{pmatrix}, \quad (2.9)$$

where ω and ξ are real parameters. The neutrino mass matrix in the mass basis is either

$$m_\nu^{\text{diag}} \approx \text{diag} \left(0, \sqrt{m_{\text{sol}}^2}, \sqrt{m_{\text{sol}}^2 + m_{\text{atm}}^2} \right), \quad (2.10)$$

for normal mass ordering or

$$m_\nu^{\text{diag}} \approx \text{diag} \left(0, \sqrt{m_{\text{atm}}^2}, \sqrt{m_{\text{sol}}^2 + m_{\text{atm}}^2} \right), \quad (2.11)$$

for the inverted one. Here, m_{sol}^2 and m_{atm}^2 are solar and atmospheric mass squared differences [56], respectively. Throughout this work we assume a mass spectrum with normal ordering.

The leptonic mixing matrix U_{PMNS} is parametrized with mixing angles θ_{ij} and phases δ_i as [56]

$$U_{\text{PMNS}} = \begin{pmatrix} e^{i\frac{\alpha_1}{2}} c_{12}c_{13} & e^{i\frac{\alpha_2}{2}} c_{13}s_{12} & e^{-i\delta} s_{13} \\ e^{i\frac{\alpha_1}{2}} (-c_{23}s_{12} - e^{i\delta} c_{12}s_{13}s_{23}) & e^{i\frac{\alpha_2}{2}} (c_{12}c_{23} - e^{i\delta} s_{12}s_{13}s_{23}) & c_{13}s_{23} \\ e^{i\frac{\alpha_1}{2}} (-e^{i\delta} c_{12}c_{23}s_{13} + s_{12}s_{23}) & e^{i\frac{\alpha_2}{2}} (-e^{i\delta} c_{23}s_{12}s_{13} - c_{12}s_{23}) & c_{13}c_{23} \end{pmatrix}, \quad (2.12)$$

with abbreviations $c_{ij} \equiv \cos \theta_{ij}$ and $s_{ij} \equiv \sin \theta_{ij}$. Considering the Majorana nature of active neutrinos in this model, there is one Dirac (δ) and 2 Majorana CP phases (α_1, α_2).

Having discussed the construction of the phenomenologically viable Yukawa matrices, we can now estimate the constraints on the parameter space from LFV processes given in table 1. The most stringent bound is coming from $\mu^+ \rightarrow e^+ \gamma$ process with the corresponding one-loop contribution approximately given by

$$\text{BR}(\mu^+ \rightarrow e^+ \gamma) \approx 2 \cdot 10^{-9} y^4, \quad (2.13)$$

for $m_N \sim \mathcal{O}(200)$ GeV and $m_\pm \sim \mathcal{O}(1)$ TeV. Here, y denotes the magnitude of the Y_{23} entries.

After comparing table 1 and eq. (2.13) we obtain $y \lesssim 0.1$. By inserting this value into eq. (2.7) we can approximately set a lower bound on the quartic coupling

$$\lambda_5 \gtrsim 4 \cdot 10^{-8}. \quad (2.14)$$

3 Dark matter production

The main goal of this section is to describe the two mechanisms for generating the abundance of keV-scale DM. First, in section 3.1, we discuss DM production via its feeble interactions with thermalized neutral (S, A) and charged scalars (σ^\pm). More precisely, the DM abundance gradually increases from the decays of heavy \mathbb{Z}_2 -odd scalars, a mechanism known as DM “freeze-in” [57] (see [58, 59] for most recent studies). We calculate the magnitude of the coupling required to obtain the observed DM energy density in the Universe.

Furthermore, DM is produced via decays of heavier right-handed neutrinos which “freeze-out” [60] from the thermal bath. The viability of such an option is explored in section 3.2. These two production mechanisms do not exclude one another and, in fact, both can significantly contribute in general scenarios.

3.1 Production via scalar decays

It is well-known that the standard thermal “freeze-out” of keV-scale DM is not feasible within the considered model.¹ Therefore, we are led toward the “freeze-in” production through which the abundance of non-thermalized N_1 is built up from the decays or annihilations of particles that are in the thermal bath. One may easily infer [33] that, in this model, the dominant production arises from the decays of neutral and charged scalars (A, S, σ^\pm) which are in the thermal bath due to the strong gauge interactions. The relevant processes are

$$A, S \rightarrow N_1 \nu_\alpha, \quad \sigma^\pm \rightarrow N_1 l_\alpha^\pm, \quad (3.1)$$

where ν_α correspond to SM neutrinos of generation α . The general form of the Boltzmann equation for the DM production via “freeze-in” reads [57]

$$\frac{dn_{N_1}}{dt} + 3H n_{N_1} = \int d\Pi_{N_1} d\Pi_\Sigma d\Pi_L (2\pi)^4 \delta^{(4)}(p_{N_1} + p_L - p_\Sigma) |\mathcal{M}|_{\Sigma \rightarrow N_1 L}^2 f_\Sigma. \quad (3.2)$$

Here, n_{N_1} is the DM number density, H is the Hubble parameter, Π_k abbreviates the phase space factor $\left(\frac{d^3 p_k}{2E_k(2\pi)^3}\right)$ for particle k , $|\mathcal{M}|_{\Sigma \rightarrow N_1 L}^2$ is the squared matrix element for the decay of a scalar into N_1 and SM lepton, and f_Σ is the distribution function of scalar particles. For brevity, we jointly denote the summation over all scalar decay channels (see eq. (3.1)) with $\Sigma \rightarrow N_1 L$.

The formula given in eq. (3.2) can be simplified to the form [57]

$$\frac{dn_{N_1}}{dt} + 3H n_{N_1} = \frac{1}{2\pi^2} \Gamma_{\Sigma \rightarrow N_1 L} m_\Sigma^2 T K_1\left(\frac{m_\Sigma}{T}\right), \quad (3.3)$$

where m_Σ is the mass of a decaying scalar, and K_1 is the modified Bessel function of the second kind.

By relating the number density of N_1 with the corresponding yield $Y = n_{N_1}/s$, where s is the entropy density, as well as using $dT = -HT dt$ we obtain the following differential equation

$$\frac{dY}{dT} = -\frac{1}{HTs} \frac{1}{2\pi^2} \Gamma_{\Sigma \rightarrow N_1 L} m_\Sigma^2 T K_1\left(\frac{m_\Sigma}{T}\right). \quad (3.4)$$

After changing the integration variable from T to $x = m_\pm/T$, as well as expanding eq. (3.4) over all Σ components, we reach the expression

$$\frac{dY}{dx} = \frac{135 M_{\text{Pl}} |y_1|^2}{1.66 \cdot 64 \pi^5 g_*^{3/2} m_\pm} x^3 \left(2 K_1(x) + r_A^3 K_1(r_A x) + r_S^3 K_1(r_S x) \right), \quad (3.5)$$

where we explicitly inserted formulae for H , s and the partial widths of the scalars [33]. For simplicity, we assumed that $y_{1\alpha}$ coupling is identical for each flavor α , therefore denoted y_1 . In eq. (3.5), g_* is the number of the degrees of freedom in the thermal bath at the time of DM production, M_{Pl} is non-reduced Planck mass, and r_A (r_S) denotes the scalar mass ratio m_A/m_\pm (m_S/m_\pm).

¹Alternatives, such as the late time entropy production [61, 62] would serve as a remedy, but then one needs to further extend the particle content of the model. Hence, we do not pursue such option.

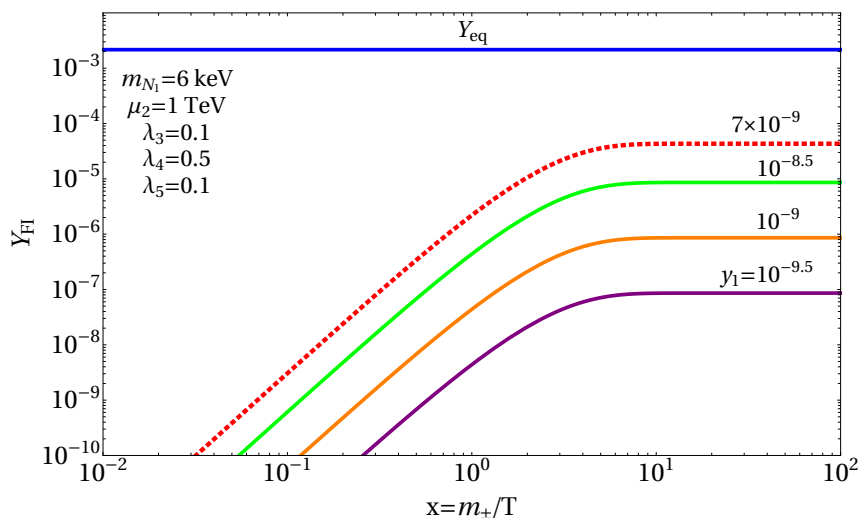


Figure 1. DM freeze-in yield as a function of $x = m_{\pm}/T$, shown for different values of y_1 . We have chosen the bare mass parameter $\mu_2 = 1000 \text{ GeV}$ and fixed the quartic terms to $\lambda_3 = \lambda_5 = 0.1$ and $\lambda_4 = 0.5$. The parameter choice of the red dashed curve leads to $\Omega h_{\text{FI}}^2 = 0.12$.

After an explicit integration where we assume the reheating temperature to be much higher than any other mass scale in our model, we obtain the following expression for the present DM yield

$$Y_{\text{FI}} = \frac{405 M_{\text{Pl}} |y_1|^2}{128\pi^4 \cdot 1.66 \cdot g_*^{3/2} m_{\pm}} \frac{2r_A r_S + r_S + r_A}{r_A r_S}. \quad (3.6)$$

By taking into account all SM and new degrees of freedom, we have $g_* \approx 114.25$ at $\mathcal{O}(10^2 - 10^3) \text{ GeV}$ temperature where the freeze-in occurs.

For the purpose of estimating the required order of magnitude for y_1 , let us further simplify eq. (3.6) by assuming $r_A = r_S = 1$ which corresponds to the $\lambda_{4,5} \rightarrow 0$ limit. We arrive at

$$Y_{\text{FI}} \approx 7.82 \times 10^{11} |y_1|^2 \left(\frac{1 \text{ TeV}}{m_{\pm}} \right). \quad (3.7)$$

By using the well-known relation between DM yield and the relic abundance

$$\Omega h_{\text{FI}}^2 = 2.742 \cdot 10^2 \left(\frac{m_{N_1}}{\text{keV}} \right) Y_{\text{FI}}, \quad (3.8)$$

we finally obtain

$$\Omega h_{\text{FI}}^2 \approx 0.12 \left(\frac{|y_1|}{2.36 \cdot 10^{-8}} \right)^2 \left(\frac{m_{N_1}}{1 \text{ keV}} \right) \left(\frac{1 \text{ TeV}}{m_{\pm}} \right). \quad (3.9)$$

In figure 1, we present the numerical solutions of eq. (3.5) for different choices of y_1 and fixed m_{N_1} . We observe that the dominant production occurs at $\mathcal{O}(1) x$. We point out that there is no strong dependence of Y_{FI} on the values of quartic couplings. Hence, any physical choice of these parameters (obeying the constraints from eq. (2.6)) yields a very similar result to the one presented in figure 1. We show the DM freeze-in for $m_{N_1} = 6 \text{ keV}$,

the value in accord with the structure formation limits given in [15]. For larger values of DM mass, coupling y_1 needs to be smaller according to eq. (3.9).

Let us finally remark that N_2 and N_3 could in principle also be produced via freeze-in from scalar decays. However, if that was the case, from eq. (3.9) we infer that the corresponding Yukawa couplings ($y_{2\alpha}, y_{3\alpha}$) would then have to be $\ll 10^{-8}$. Such small couplings would not be sufficiently large to generate eV-scale neutrino masses (see eq. (2.7)). The required strength of ($y_{2\alpha}, y_{3\alpha}$) puts N_2 and N_3 in thermal equilibrium with SM particles.

3.2 Production via N_2 decays

As already discussed in section 2, the \mathbb{Z}_2 -odd sector consists in total of three right-handed neutrinos, neutral (A and S) and charged scalars σ^\pm . The mass spectrum is assumed to be $m_{N_1} \ll m_{N_2} \lesssim m_{N_3} < (m_\pm, m_A, m_S)$. This spectrum has several appealing features. Since scalars are heavier than right-handed neutrinos, potential breaking of parity symmetry induced by renormalization group evolution is evaded [63]. Additionally, models with a low scale DM, in comparison with conventional $\mathcal{O}(10^2)$ GeV DM, manifest improvement in the predictions for the small scale structure [64].

Due to small Yukawa couplings of N_1 ($y_{1\alpha} \ll y_{2\alpha}, y_{3\alpha}$) (see section 3.1), the decay rate for the process $N_2 \rightarrow N_1 L_\alpha \bar{L}_\beta$ is much smaller in comparison to the processes involving heavier \mathbb{Z}_2 -odd particles. We have numerically checked that such decays occur only after N_2 undergoes a successful thermal freeze-out. Therefore, the amount of DM produced from N_2 decays is determined by the freeze-out abundance of N_2 , which we calculate in the following.

The Boltzmann equation for each of the considered \mathbb{Z}_2 -odd particles ($N_2, N_3, \sigma^\pm, A, S$, commonly denoted as $\chi_i, i = 1 \dots N$) is given as [65]

$$\frac{dn_i}{dt} + 3Hn_i = - \sum_{j=1}^N \langle \sigma_{ij} v_{ij} \rangle (n_i n_j - n_i^{\text{eq}} n_j^{\text{eq}}) \quad (3.10)$$

$$- \sum_{j \neq 1}^N \left(\langle \sigma'_{Xij} v_{ij} \rangle (n_i n_X - n_i^{\text{eq}} n_X^{\text{eq}}) - \langle \sigma'_{Xji} v_{ij} \rangle (n_j n_X - n_j^{\text{eq}} n_X^{\text{eq}}) \right) \quad (3.11)$$

$$- \sum_{j \neq i}^N \left(\Gamma_{ij} (n_i - n_i^{\text{eq}}) - \Gamma_{ji} (n_j - n_j^{\text{eq}}) \right), \quad (3.12)$$

where the terms in the first, second and third row on the right hand side represent (co)annihilations, scatterings and decays, respectively. Here, $\sigma_{ij} = \sum_X \sigma(\chi_i \chi_j \rightarrow X)$ is the (co)annihilation cross section for \mathbb{Z}_2 -odd particles into SM particles, $\sigma'_{Xij} = \sum_Y \sigma(\chi_i X \rightarrow \chi_j Y)$ is the cross section for scattering with SM bath (denoted with X and Y), and $\Gamma_{ij} = \sum_X \Gamma(\chi_i \rightarrow \chi_j X)$ are decay rates.

Since all heavier \mathbb{Z}_2 particles eventually decay into N_2 , we may define the total N_2 number density as²

$$n = \sum_i n_i. \quad (3.13)$$

Considering the size of the coupling constants in our setup, all the χ_i are in the thermal equilibrium at $\mathcal{O}(1)$ TeV temperatures with the number density equal to

$$n_i^{\text{eq}} = \frac{T}{2\pi^2} \sum_i g_i m_i^2 K_2\left(\frac{m_i}{T}\right), \quad (3.14)$$

where g_i denotes the number of internal degrees of freedom of species i . By employing eqs. (3.12) and (3.13) and as well as the relation [65]

$$\frac{n_i}{n} \simeq \frac{n_i^{\text{eq}}}{n^{\text{eq}}}, \quad (3.15)$$

we arrive at the standard Boltzmann equation for the evolution of N_2 number density [66]

$$\frac{dn}{dt} = -3Hn - \langle\sigma_{\text{eff}}v\rangle(n^2 - n_{\text{eq}}^2), \quad (3.16)$$

where $\langle\sigma_{\text{eff}}v\rangle$ is given by [65]

$$\langle\sigma_{\text{eff}}v\rangle = \sum_{i,j} \langle\sigma_{ij}v_{ij}\rangle \frac{n_i^{\text{eq}} n_j^{\text{eq}}}{(n^{\text{eq}})^2}. \quad (3.17)$$

Note that in eq. (3.16) there are no scattering terms as well as decays. After employing the summation given in eq. (3.13) these contributions get removed as they do not change the total number of \mathbb{Z}_2 -odd particles.

Changing variables to dimensionless quantities $Y = n/s$ and $y = m_{N_2}/T$ yields

$$\frac{dY}{dy} = \sqrt{\frac{\pi g_*}{45}} \frac{M_{\text{Pl}} m_{N_2}}{y^2} \langle\sigma_{\text{eff}}v\rangle (Y_{\text{eq}}^2 - Y^2). \quad (3.18)$$

We have evaluated eq. (3.18) with `micrOMEGAs` [67] and compared with the output of `MadDM` [68]. The Yukawa couplings of N_2 and N_3 are generated as described in section 2.2. Therefore, in the code, only the points in the parameter space that are fully consistent with the neutrino mixing angles and mass squared differences are evaluated.

In figure 2, we show the time evolution of Y for several different couplings and right-handed neutrino masses. As can be seen, the yield can reach high values for large λ_5 . This is because, in such case, the Yukawa couplings $y_{2\alpha}$ are small (see eq. (2.7)) and N_2 does not remain in thermal equilibrium for long (freeze-out takes place at higher temperatures with respect to the usual $x \approx 20$).

It is important to point out the coannihilation processes which are dominant if the scalars and N_2 have similar masses. For $m_{\pm} > m_{N_2} \gtrsim 0.8 m_{\pm}$, Y is generally lowered by several orders of magnitude with respect to the general case with hierarchical fermion and scalar masses. This can also be seen in figure 2 where the green (purple) curve indicates the case with $m_{N_2} = 0.85 m_{\pm}$ ($m_{N_2} = 0.95 m_{\pm}$). We will particularly emphasize the importance

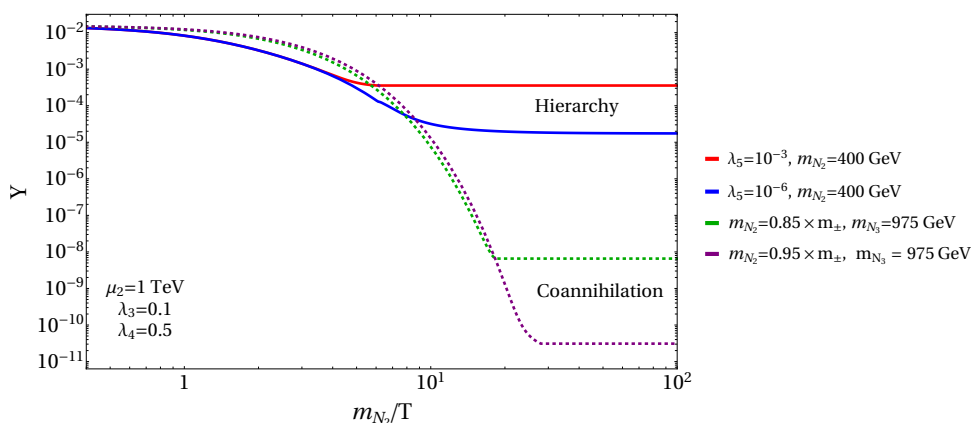


Figure 2. Evolution of the total N_2 yield for different values of λ_5 in a scenario with vastly different scalar and N_2 masses (red and blue curve) as well as in the case when the masses are similar (green and purple dashed curve). In the latter case, due to larger interaction strength, N_2 remains in thermal equilibrium longer, yielding orders of magnitude smaller N_2 abundance with respect to the former case.

of coannihilations when exploring viable joint parameter space for DM and leptogenesis (see section 5).

The contribution to the DM relic abundance from N_2 decays is

$$\Omega h_{N_2 \rightarrow N_1}^2 = \frac{m_{N_1}}{m_{N_2}} \Omega h_{N_2}^2, \quad (3.19)$$

where the relation between $\Omega h_{N_2}^2$ and Y matches the one in eq. (3.8).

Taking into account the complementary contribution from scalar decays given in section 3.1, the requirement for having the amount of DM in accordance with the measurements [69] is

$$\Omega h_{N_2 \rightarrow N_1}^2 + \Omega h_{\text{FI}}^2 = 0.12. \quad (3.20)$$

As already discussed, we construct the Yukawa couplings of N_2 and N_3 in agreement with the results from neutrino oscillation experiments. Such a choice already unambiguously fixes $\Omega h_{N_2 \rightarrow N_1}^2$ contribution, meaning that if it is already overshooting the observed value of 0.12, the corresponding parameter choice is excluded. Otherwise, the Yukawa couplings of N_1 could be accommodated in such a way to satisfy eq. (3.20). Our findings in section 3.1 indicate $y_{1\alpha} \lesssim 10^{-8}$.

In figure 3, results for $\Omega h_{N_2 \rightarrow N_1}^2$ are shown as a function of m_{N_2} . For $m_{N_2} \lesssim 500$ GeV one may infer that in the case of larger λ_5 couplings, $\Omega h_{N_2 \rightarrow N_1}^2$ poses the dominant contribution to DM relic density since, due to the weakness of $N_2 N_2 \rightarrow \text{SM SM}$ channels, N_2 undergoes freeze-out very early. On the other hand, for higher m_{N_2} , additional annihilation channels involving new scalars become dominant, increasing the overall cross section, and correspondingly reducing $\Omega h_{N_2 \rightarrow N_1}^2$. In contrast, for very tiny λ_5 (pink curve), N_2 annihilates efficiently even in the absence of coannihilation processes, and therefore the

²The procedure also holds for almost degenerate N_2 and N_3 (see section 4).

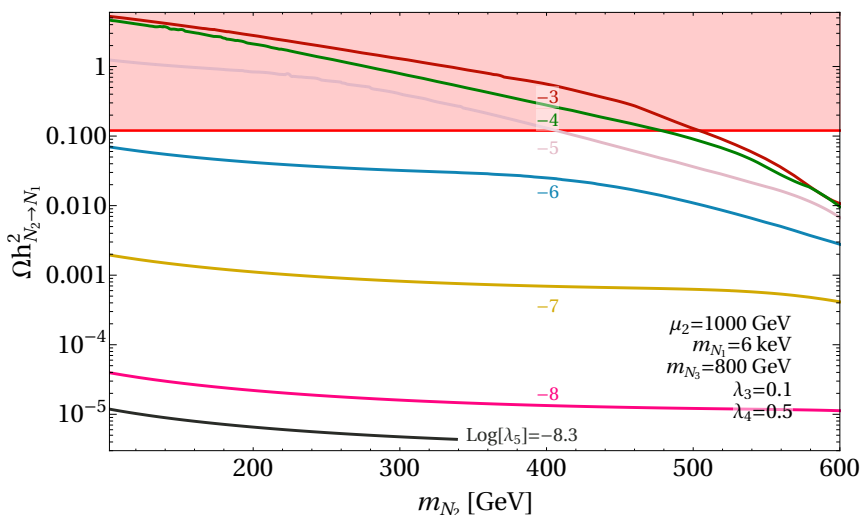


Figure 3. $\Omega h^2_{N_2 \rightarrow N_1}$ for different choices of the coupling constant λ_5 . In general, larger values of this coupling lead to larger DM production from N_2 decays. Every parameter point in the red shaded region is overproducing DM and is thus ruled out. The truncation of the $\lambda_5 = 10^{-8.3}$ curve is due to the bounds from LFV processes.

m_{N_2} dependence flattens. We stress here again that λ_5 can not be pushed to arbitrarily low values, because of the constraints from lepton flavor violation experiments (see section 2). In conclusion, two different behaviors can be seen in the light of eq. (3.20). For small λ_5 , the parameters $y_{1\alpha}$ and m_{N_2} are practically independent of each other, whereas a cutoff value of m_{N_2} below which DM would be overproduced is reached for larger λ_5 .

4 Leptogenesis

In order to generate the observed BAU we study the option where initially an asymmetry in the lepton sector is produced [70]. The production of lepton asymmetry is dubbed leptogenesis. In the simplest realizations, the models featuring leptogenesis incorporate the seesaw type-I mechanism and rely on CP violating decays of hierarchical heavy right-handed neutrinos. The produced lepton asymmetry can be partially converted to the baryon asymmetry due to the existence of non-perturbative processes in thermal equilibrium [71]. These processes, called sphalerons, violate $B+L$, but conserve $B-L$ numbers and thus allow for an asymmetry conversion between the two sectors. When the temperature drops below $T_s = 131.7 \text{ GeV}$ [72], the sphalerons decouple from the thermal bath and the asymmetry conversion ceases.

The drawback of the above mentioned realization is that in order to generate the observed BAU the masses of the right-handed neutrinos have to be at least of $\mathcal{O}(10^8) \text{ GeV}$ [73, 74] which is not experimentally reachable.³ It is, however, possible to lower the needed mass scale below TeV. In this paper we present such an option in the

³Taking into account flavor effects [75–78], recent studies suggest that masses as low as 10^6 GeV are consistent with successful thermal leptogenesis, which are nevertheless difficult to access experimentally [79].

frame of the Scotogenic model. We rely on two competing mechanisms to create a sufficient baryon asymmetry.

Firstly, we incorporate oscillations among the right-handed neutrinos (ARS) [2, 80–83] which transfer the asymmetry to the SM leptons via lepton portal (see eq. (2.4)). Secondly, we also take into account decays of the new Σ doublet which at finite temperatures serves as an additional source of CP violation. This process is suppressed by a mass insertion factor $(m_{N_2}/T)^2$ and was neglected in the past literature. However, recently it has been shown that scalar decays are important [3] and can even dominate ARS in some regions of the parameter space [84]. Hence, we take both ARS and scalar decays into account.

It is worth noting that several attempts to implement leptogenesis in the Scotogenic model had already been made (see e.g. [85, 86]). In refs. [87, 88], the authors were able to achieve a low scale leptogenesis without imposing mass degeneracies between right-handed neutrinos.

For studying leptogenesis, we apply the procedure presented for seesaw type-I mechanism in ref. [84] to the Scotogenic model. In other words, we replace SM Higgs doublet with Σ in the lepton portal term. Only the two right-handed neutrinos, namely N_2 and N_3 participate in the lepton asymmetry production. In the density matrix formalism, the production of a lepton asymmetry via both ARS mechanism and scalar decays is automatically included. Adopting the notation from [84], we work with yields ρ associated with the number density matrices of right-handed neutrinos

$$\rho = \frac{n}{s}, \quad \bar{\rho} = \frac{\bar{n}}{s}. \quad (4.1)$$

Here, n is a 2×2 matrix with diagonal elements associated to the densities and the off-diagonal entries parametrizing the mixing between N_2 and N_3 fields. The states with the opposite helicity are indicated with the overline. In order to calculate the lepton asymmetry δY_l , for each flavor l , we solve the following set of coupled differential equations

$$\begin{aligned} \frac{d\rho_{\alpha\beta}}{dz} = & -\frac{i}{zH} [\mathcal{E}_N, \rho]_{\alpha\beta} - \frac{1}{2zHs} \left\{ \gamma^{LC} + \gamma^{LV}, \frac{\rho}{\rho_{\text{eq}}^N} - \mathbb{1} \right\}_{\alpha\beta} \\ & + \frac{\delta Y_l}{2n_{\text{eq}}^L zH} \left((\gamma_{WQ,l}^{LC} - \gamma_{WQ,l}^{LV}) + \frac{1}{2} \left\{ \gamma_{WC,l}^{LC} - \gamma_{WC,l}^{LV}, \frac{\rho}{\rho_{\text{eq}}^N} \right\} \right)_{\alpha\beta}, \end{aligned} \quad (4.2)$$

$$\begin{aligned} \frac{d\bar{\rho}_{\alpha\beta}}{dz} = & -\frac{i}{zH} [\mathcal{E}_N, \bar{\rho}]_{\alpha\beta} - \frac{1}{2zHs} \left\{ \gamma^{LC*} + \gamma^{LV*}, \frac{\bar{\rho}}{\rho_{\text{eq}}^N} - \mathbb{1} \right\}_{\alpha\beta} \\ & - \frac{\delta Y_l}{2n_{\text{eq}}^L zH} \left((\gamma_{WQ,l}^{LC*} - \gamma_{WQ,l}^{LV*}) + \frac{1}{2} \left\{ \gamma_{WC,l}^{LC*} - \gamma_{WC,l}^{LV*}, \frac{\bar{\rho}}{\rho_{\text{eq}}^N} \right\} \right)_{\alpha\beta}, \end{aligned} \quad (4.3)$$

$$\begin{aligned} \frac{d\delta Y_l}{dz} = & \frac{1}{zHs\rho_{\text{eq}}^N} \text{Tr} \left\{ (\gamma_l^{LC} - \gamma_l^{LV}) \rho \right\} \\ & - \frac{1}{zHs\rho_{\text{eq}}^N} \text{Tr} \left\{ (\gamma_l^{LC*} - \gamma_l^{LV*}) \bar{\rho} \right\} - \frac{\delta Y_l}{zHn_{\text{eq}}^L} \text{Tr} \left\{ \gamma_{WQ,l}^{LC} + \gamma_{WQ,l}^{LV} \right\} \\ & - \frac{\delta Y_l}{2zHn_{\text{eq}}^L} \frac{1}{\rho_{\text{eq}}^N} \text{Tr} \left\{ \rho (\gamma_{WC,l}^{LC} + \gamma_{WC,l}^{LV}) \right\} - \frac{\delta Y_l}{2zHn_{\text{eq}}^L} \frac{1}{\rho_{\text{eq}}^N} \text{Tr} \left\{ \bar{\rho} (\gamma_{WC,l}^{LC*} + \gamma_{WC,l}^{LV*}) \right\}, \end{aligned} \quad (4.4)$$

where the parameter z equals $z = T_s/T$ and $T_s \sim 131.7 \text{ GeV}$ is the temperature at which sphalerons decouple. Integration is performed between $z_{\text{start}} = T_{\text{sph}}/T = 10^{-3}$ and $z_{\text{stop}} = 1$ with the initial conditions $\rho_{\alpha\beta}(z_{\text{start}}) = 0$, $\bar{\rho}_{\alpha\beta}(z_{\text{start}}) = 0$, $\delta Y_l(z_{\text{start}}) = 0$, where indices α and β run from 1 to 2 due to the relevance of only two right-handed states. In order to obtain the final baryon asymmetry δY_B we evaluate $-\frac{2}{3} \sum_{l=e,\mu,\tau} \delta Y_l$ [89].

The most general form for the diagonal matrix elements of \mathcal{E}_N is

$$\mathcal{E}_{N,\alpha\alpha} \equiv \frac{1}{n_{\text{eq}}^N} \int_0^\infty \frac{dk k^2}{2\pi^2} \frac{\sqrt{k^2 + m_{N_\alpha}^2}}{\exp\left[\frac{\sqrt{k^2 + m_{N_\alpha}^2} z/T_s\right] + 1}, \quad (4.5)$$

with

$$n_{\text{eq}}^N = \frac{g}{2\pi^2} \int_0^\infty dp \frac{p^2}{\exp\left[\frac{\sqrt{p^2 + m^2} z/T_s\right] + 1}, \quad \rho_{\text{eq}}^N = \frac{n_{\text{eq}}^N}{s}, \quad (4.6)$$

where $g = 1$ for right-handed neutrinos. In the relativistic case, the expression in eq. (4.5) reduces to [84]

$$\mathcal{E}_N \approx \frac{0.46 z}{2T_s} m_{N_2}^2 \begin{pmatrix} 1 & 0 \\ 0 & (1 + \delta_M)^2 \end{pmatrix}, \quad (4.7)$$

where we introduced the level of degeneracy between the masses of N_2 and N_3 defined via relation $m_{N_3} = m_{N_2}(1 + \delta_M)$. Note that, as we will show later explicitly, $\delta_M \lesssim 10^{-8}$ for a successful leptogenesis [84].

In the numerical evaluation, we have employed the general expressions for all the terms involving right-handed neutrinos given in eqs. (4.2), (4.3) and (4.4), i.e. we did not take the relativistic approximation. This is because our findings indicate that successful leptogenesis occurs when $m_{N_{2,3}} \gtrsim T_s$, for which the relativistic approximation does not hold. In contrast, the equilibrium number density for SM leptons is evaluated in the relativistic approximation

$$n_{\text{eq}}^L = \frac{3\zeta(3)}{2\pi^2} \left(\frac{T_s}{z}\right)^3, \quad (4.8)$$

since the heaviest SM lepton is two orders of magnitude lighter with respect to T_s . In eqs. (4.2), (4.3) and (4.4) we denote the reaction densities for ARS and Σ decays with γ^{LC} and γ^{LV} , respectively, where LC and LV are abbreviations for the lepton number conserving and violating processes. Washout terms are labeled with extra subscripts - WC or WQ where the former (latter) indicates lepton asymmetry loss due to classical (quantum) effects. The derivation of the reaction densities and the washout terms is outlined in appendix A. Here, we would like to emphasize the scalar mass treatment in such calculation. First of all, in evaluating the integrals over the phase space, for simplicity and the possibility of analytical evaluation we take $m_\pm = m_S = m_A$. We have nevertheless checked numerically that introducing the splitting between these masses, arising from $\lambda_{4,5} \neq 0$ (see eq. (2.2)), yields an insignificant change of reaction densities and washout terms which does not influence the final value of baryon asymmetry.

In addition to the bare masses for new scalars, which are $\mathcal{O}(10^2-10^3)$ GeV, the thermal corrections are important. We therefore solve eqs. (4.2), (4.3) and (4.4) in two different temperature regimes. For the low temperatures, we neglect thermal corrections and use

$$m_{\pm}^2 = \mu_2^2 + \lambda_3 v^2, \quad (4.9)$$

while for high temperatures we adopted the expression for thermal corrections, calculated in refs. [90] and [91]

$$m_{\pm}^2 = \left(\frac{3\lambda_2 + 2\lambda_3 + \lambda_4}{12} + \frac{3g^2 + g'^2}{16} \right) \frac{T_s^2}{z^2}, \quad (4.10)$$

where g and g' are $SU(2)_L$ and $U(1)_Y$ coupling constants, respectively. In total, we are using the following prescription for the scalar mass

$$m_{\pm}^2 = \begin{cases} \left(\frac{3\lambda_2 + 2\lambda_3 + \lambda_4}{12} + \frac{3g^2 + g'^2}{16} \right) \frac{T_s^2}{z^2} + \mu_2^2 + \lambda_3 v^2, & \text{if } \left(\frac{3\lambda_2 + 2\lambda_3 + \lambda_4}{12} + \frac{3g^2 + g'^2}{16} \right) \frac{T_s^2}{z^2} \geq 3(\mu_2^2 + \lambda_3 v^2), \\ \mu_2^2 + \lambda_3 v^2, & \text{otherwise.} \end{cases} \quad (4.11)$$

The boundary temperature separating the two regimes is determined by the condition that the thermal mass is equal to 3 times the bare mass. We checked numerically that varying the boundary temperature by $\mathcal{O}(1)$ numbers does not affect the final results. Below this boundary we drop thermal corrections to the leptons, considering them to be effectively massless at low temperatures. Let us note that in both regimes, the phase space for Σ decays is kinematically open, i.e. scalar masses are always larger than the total mass of decay products (right-handed neutrino and SM lepton). We would also like to stress that we have checked that in the relativistic limit our results match those presented in ref. [84].

The temperature dependence of the γ terms (reaction densities and washout terms) is shown in figure 4 where we compare the relativistic regime (neglecting all bare masses and m_N , solid lines), with the general aforementioned approach (taking the scalar mass as given in eq. (4.11) and accounting for the effects from non-zero right-handed neutrino masses, dashed lines). We note the suppression of γ^{LV} terms at low temperatures. This is because, in the absence of thermal effects, it is much less probable to have an on-shell mediator particle [3]. The γ^{LC} factors feature an opposite effect - get larger at $z \sim 0.1$ where the phase space for scattering processes between L and Σ increases. This is because m_{\pm} is fixed to the bare mass and the lepton mass is gradually decreasing due to the dropping temperature. We also observe that all γ terms in both panels become Boltzmann suppressed at $z \sim 1$.

Radiative generation of neutrino masses introduces a suppression factor of $\lambda_5^2/16\pi^2$ which lifts up the overall strength of $(y_{2\alpha}, y_{3\alpha})$ in comparison to the corresponding values in seesaw type-I model. Numerically, these couplings can not be made smaller than $\mathcal{O}(10^{-6})$ in our model.

Such interaction strength may pose a problem for generating observed BAU due to strong washout of the generated lepton asymmetry at late times. Let us note that, in principle, leptogenesis can be successfully achieved with such large couplings [84] when ξ (see eq. (2.9)) is large. In addition, as pointed out in ref. [83], the large imaginary

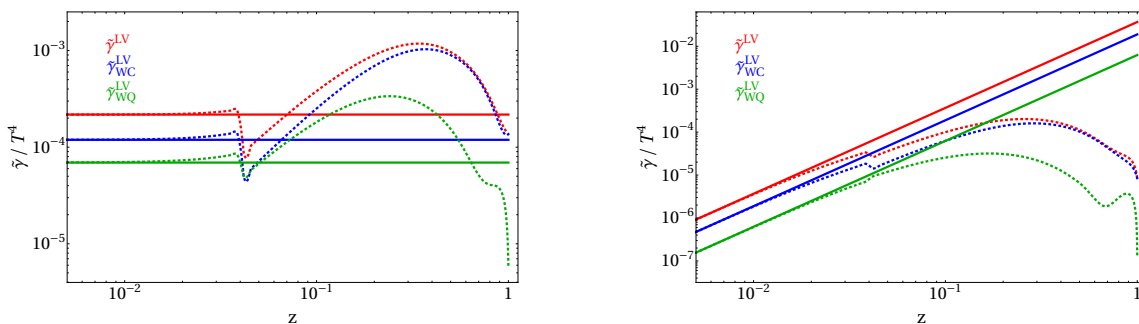


Figure 4. Temperature dependence of γ^{LC} and γ^{LV} (see appendix A) with stripped-off Yukawa couplings (denoted with tilde). We take $\mu_2 = 1$ TeV, $m_{N_2} = 400$ GeV and the values of the scalar couplings are given in table 2. The T^4 dependence is factored out in both panels. In both panels, the dashed (solid) lines correspond to the regime where bare (thermal) mass contributes. At $z \approx 0.05$ thermal corrections are as big as the bare mass, indicating a change of considered regime.

quartic coupling	value
λ_2	0.3
λ_3	-0.27
λ_4	2
λ_5	2

Table 2. Values of the scalar couplings motivated by a maximal reduction of the strength of $y_{2,3\alpha}$. Values are in accord with the limits coming from the requirement for the stability of the Higgs potential (see section 2.1), as well as perturbativity.

entries (induced by ξ) may lead to certain cancellations between different terms in the Boltzmann equations, preventing the generated lepton asymmetry from strong washout effects. However, this reasoning can not be applied to our model because pushing ξ to large values drastically increases the strength of washout effects. In fact, even the asymmetry produced at very late times will be washed out efficiently before sphaleron decoupling.

In order to reduce the Yukawa couplings as much as possible we choose the quartic couplings as given in table 2, consistent with the limits from the stability of the Higgs potential (see section 2.1). With couplings chosen in such a way we find that successful leptogenesis may only be generated with $m_{N_2} \gtrsim \mathcal{O}(10^2)$ GeV. For the right-handed neutrino masses smaller than T_s , the washout effects at $z \gtrsim 0.1$ remove the vast majority of the generated asymmetry. In contrast, if $m_{N_2} \gtrsim T_s$, the right-handed neutrino abundance becomes Boltzmann suppressed at $z \gtrsim 10^{-1}$ which strongly suppresses washout integrals given in appendix A. While at such late times the asymmetry production is suppressed in a similar way, we find that the strong production at $z \ll 1$ suffices for generating the observed $\delta Y_B = 0.86 \times 10^{-10}$ [69] baryon asymmetry. Throughout the paper, without the loss of generality, we use only positive values of ξ . This parameter enters exponentially ($\exp^{\pm\xi}$) in the Yukawa matrix and thus its negative values would not change qualitatively the overall picture for BAU generation.

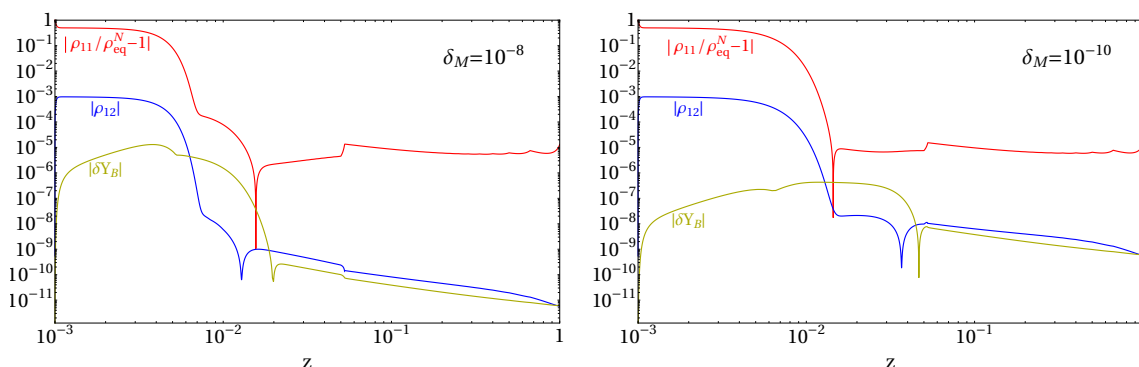


Figure 5. Evolution of $\rho_{11}(z)$, $\rho_{12}(z)$ and $\delta Y_B(z)$ for two different values of δ_M . We have fixed $\omega = \pi/4$, $\delta = -\pi/2$, $\xi = 2$ and $\alpha_1 = \alpha_2 = 0$. The asymmetry for $\delta_M = 10^{-8}$ is produced earlier but also more strongly washed out with respect to $\delta_M = 10^{-10}$ case. We found no significant dependence of the generated asymmetry on the values of Dirac (δ) and Majorana (α_1 and α_2) CP phases.

In figure 5 we show the deviation from the equilibrium value of N_2 number density ($\rho_{11}(z)/\rho_{\text{eq}}^N - 1$) and the mixing between N_2 and N_3 ($\rho_{12}(z)$). The evolution of the baryon asymmetry δY_B is also presented for $m_{N_2} = 200$ GeV, $\mu_2 = 800$ GeV and with the following values of angles and phases $\omega = \pi/4$, $\delta = -\pi/2$, $\alpha_1 = \alpha_2 = 0$. The kink at $z \approx 0.05$ is due to the change of regimes for the thermal masses (see eq. (4.11)). The narrow feature at slightly smaller temperatures indicates the sign change of δY_B . A significant deviation of ρ_{11} and ρ_{22} from ρ_{eq}^N as well as avoiding very tiny off-diagonal matrix elements (ρ_{12} and ρ_{21}) are crucial for the successful leptogenesis. This is only ensured for $m_{N_2} \gtrsim T_s$, as discussed above. We would like to stress that there is no significant dependence of the generated lepton asymmetry on α_1 and α_2 . In addition, note that δ_M has to be tiny for generating the asymmetry, implying strong level of degeneracy between heavier right-handed neutrino masses. From the shape of $\rho_{12}(z)$ and $\delta Y_B(z)$ curves for $\delta_M = 10^{-8}$ (left panel) and $\delta_M = 10^{-10}$ (right panel) we infer that in the former case, the washout effects are effective throughout a longer time period. This can be understood from the temperature scale z_{osc} at which the oscillation among right-handed neutrinos is most effective. It is given by [2]

$$z_{\text{osc}} = T_s \left(2\sqrt{45}/(4\pi^3 g_*) m_{N_2} \delta_M M_{\text{Pl}} \right)^{-1/3}. \quad (4.12)$$

For smaller δ_M , the asymmetry is produced later which leads to a weaker washout and consequently larger δY_B .

In figures 6a and 6b we show the generated baryon asymmetry in $\xi - M_{N_2}$ plane for two fixed values of λ_5 . From this figure one can also qualitatively infer the dependence of the generated baryon asymmetry on the Yukawa interaction strength since different values of λ_5 imply different $y_{2,3\alpha}$. As expected, there is only a mild dependence on m_{N_2} which does not alter the Yukawa strength dramatically. From figure 6c one can deduce that there is a saturation effect for $\delta Y_B \gtrsim 4 \times 10^{-9}$, where the dependence of the asymmetry on λ_5 flattens. Furthermore, one should note that in case of large λ_5 , δY_B is generally very weakly dependent on this quartic coupling. Finally, figure 6d summarizes the interplay between λ_5

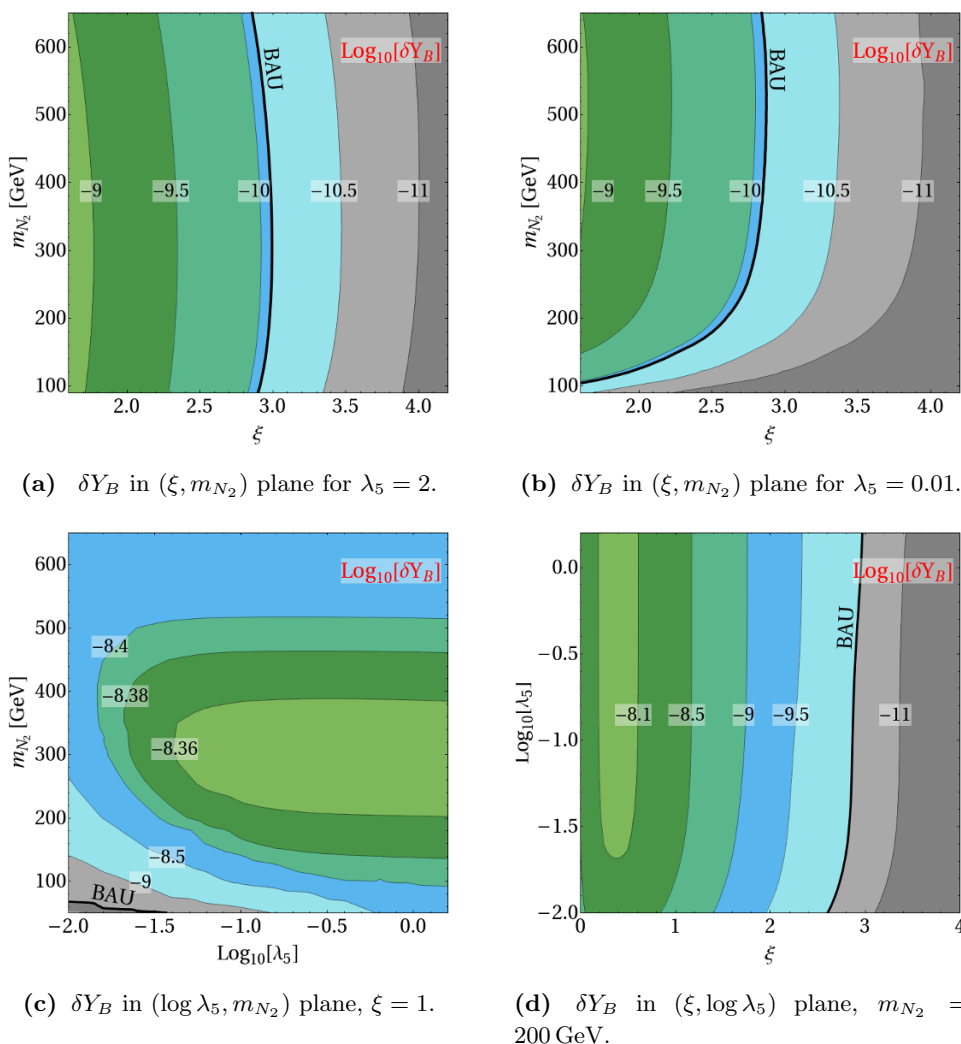


Figure 6. δY_B shown for various choices of parameters. The level of degeneracy between N_2 and N_3 masses is fixed to $\delta_M = 10^{-10}$ in all panels. The solid black line indicates the value of the observed baryon asymmetry of the Universe.

and ξ . Intuitively, it is clear that the largest asymmetry production happens for low ξ and high λ_5 values, whereas going to the opposite regime favors washout, hence leading to a reduced final asymmetry. Interestingly, ξ has a stronger impact than the quartic coupling, despite the fact that in the considered parameter range both quantities vary the size of the Yukawa coupling by almost two orders of magnitude. The reason for this behavior is that while λ_5 just sets the overall strength of the Yukawa couplings, ξ enters in the Yukawa matrix in a specific pattern, leading to rather non trivial effects.

From figure 7 we observe that when considering only ARS, $|\delta Y_B|$ decreases roughly by one order of magnitude with respect to the values from figure 6a. What can also be inferred is the stronger dependence on m_{N_2} in absence of scalar decays.

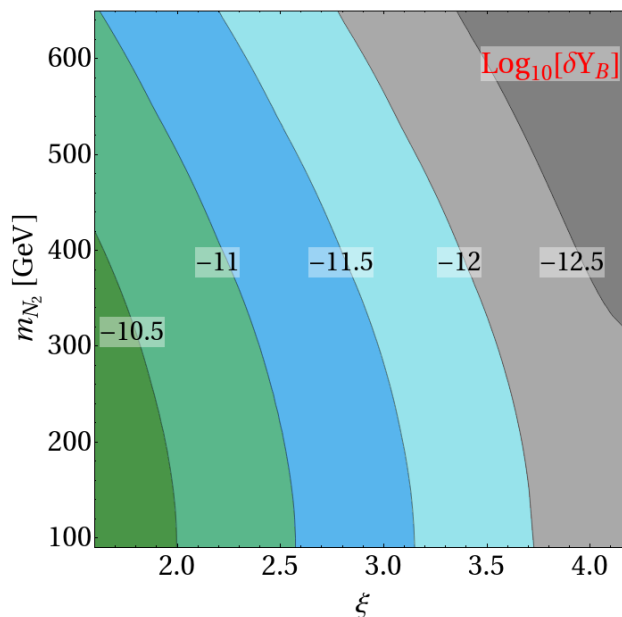


Figure 7. Same as figure 6a but with only ARS contribution taken into account. This generically leads to a smaller final asymmetry production as can also be seen in figure 8.

We find that it is very important to account for the asymmetry production via new scalar decays in addition to ARS. Actually, turning either of the two production mechanisms off yields a loss of a substantial amount of the generated baryon asymmetry, as evident from figure 8. Obviously, there is a strong interplay between both regimes, because their individual contributions are rather small and do not trivially add up to yield the final result. Each of the regimes individually leads to a similar final asymmetry contribution. Initially, the asymmetry generation is governed exclusively by the ARS mechanism and Σ decays start to compete at $z \approx 0.005$. Interestingly, the intermediate asymmetry produced by Σ decays is one order of magnitude larger compared to the one in ARS, but tends to get washed out more strongly. At $z \approx 0.05$, δY_B changes its sign and later fades away due to washout effects. At even smaller temperatures all regimes feature similar asymmetry evolution (governed by washout effects) until sphaleron processes eventually decouple. Again, including both regimes does not translate into summing up their individual contributions. The asymmetry generation in the presence of both mechanisms shows a qualitatively different behavior, featuring a peak at larger temperatures. Non-linear combination of ARS and scalar decays explicitly demonstrates the absence of the weak-washout regime in which the contributions would add trivially.

5 Identifying the viable joint parameter space for DM and leptogenesis

In sections 3 and 4 we separately scrutinized DM and BAU production. Therefore, the question arises whether it is possible to find regions in the parameter space for which the observed values of DM relic abundance and $|\delta Y_B|$ are simultaneously reached. Generally, these two mechanisms have conflicting requirements on the strength of the Yukawa coupling. In

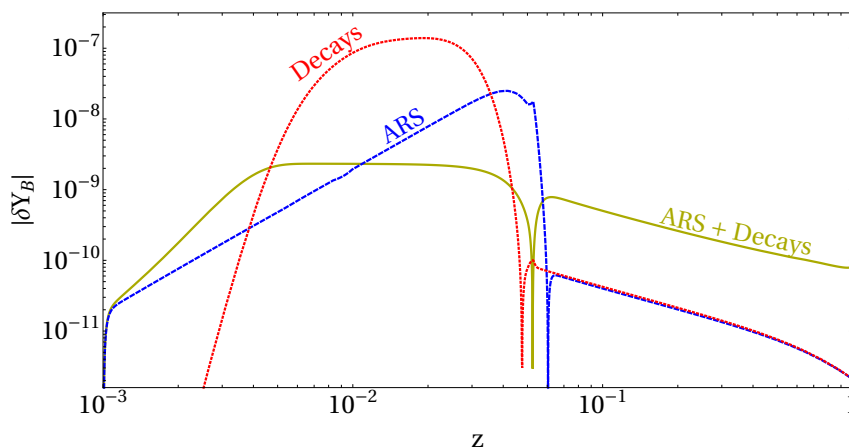


Figure 8. Time evolution of $|\delta Y_B|$ for the cases of Σ decays (red dotted) and ARS (blue dashed) shown together with the general case when the both production mechanisms are included (yellow solid). We take $\xi = 3$ and $m_{N_2} = 400$ GeV. It is evident that the combination of both production mechanisms suffices to account for BAU, in contrast with the individual contributions.

order not to overproduce DM from decays of N_2 (section 3.2), sufficiently large Yukawa interactions are required. On the other hand, leptogenesis relies on weak interactions as otherwise the washout effects would easily destroy any generated lepton asymmetry.

A key method for curing this contrast is to impose coannihilations between the right-handed neutrinos and the scalars (especially the lightest scalar σ^\pm) by choosing them to have similar masses. The coannihilation case was already discussed in section 3 (see in particular figure 2). Such a regime opens up new scalar annihilation channels which do not rely on the strength of the second and third generation Yukawa couplings. Therefore, a huge suppression of the relic density can be achieved for Yukawa couplings set low enough to generate a significant lepton asymmetry.

We now revisit BBN constraints on the N_2 decays introduced in section 2.1. By requiring N_2 to decay before $t_{\text{BBN}} \sim 1$ sec we obtain

$$|y_{2\alpha}|^2 \gtrsim 6.3 \cdot 10^{-7} \left(\frac{m_\pm}{1 \text{ TeV}}\right)^4 \left(\frac{1 \text{ TeV}}{m_{N_2}}\right)^5 \left(\frac{10^{-8}}{|y_1|}\right)^2. \quad (5.1)$$

After a more careful analysis of the processes influencing the primordial abundances of light elements this limit gets significantly relaxed. Following ref. [32], we infer that for $Y m_{N_2} \lesssim \mathcal{O}(10^{-9})$, where Y is the DM yield (see eq. (3.18)), the BBN limit given in eq. (5.1) is relaxed by roughly three orders of magnitude. This is because, in our model, N_2 has only leptonic decays and the strongest effect from charged leptons on the primordial abundances of the light nuclei is coming from the photodissociation process which is most effective at approximately 1000 seconds after the Big Bang.

We calculated the relic density $\Omega h_{N_2}^2$ for different choices of μ_2 and also evaluated δY_B in a range of m_{N_2} and ξ , taking into account $m_{N_2} < m_\pm$. Then, we determined which region in the parameter space is ruled out due to low $y_{2\alpha}$. We conservatively adopted the

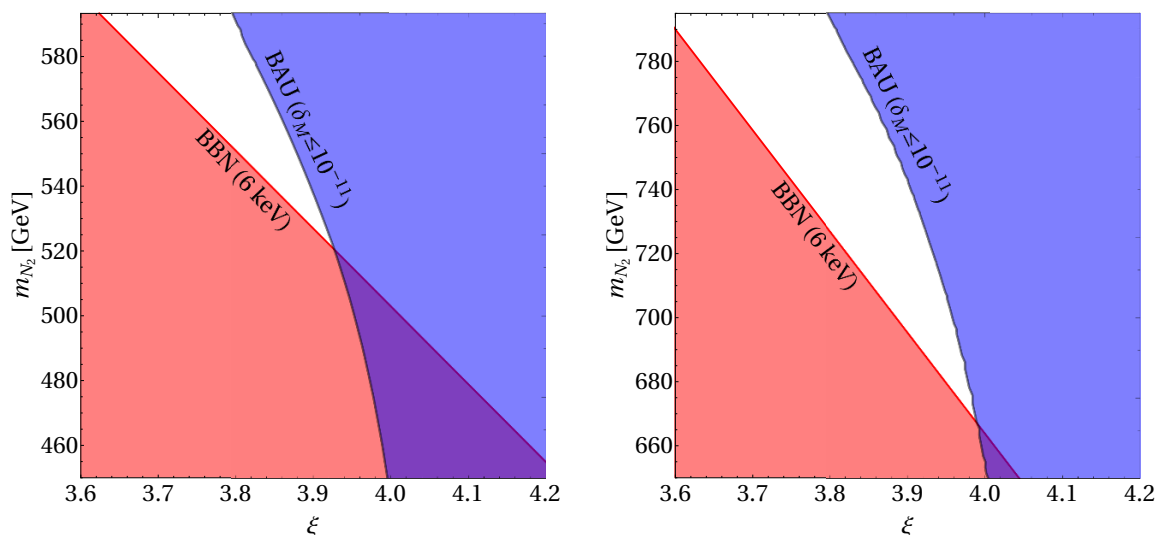


Figure 9. Allowed region in the (ξ, m_{N_2}) parameter space for fixed $\mu_2 = 600$ GeV (left) and $\mu_2 = 800$ GeV (right). The BBN exclusion limits are shown in red, while the blue shaded region does not produce a sufficiently large baryon asymmetry. We observe that there is a substantial region consistent with BBN limits in which the correct amounts of DM and baryon asymmetry can be obtained.

limit for the most stringent BBN constraint on pure leptonic decays ($\tau^+\tau^-$, see ref. [32]). In figure 9 we present the results for two different choices of μ_2 . In both panels, the blue region is indicating excluded parameter space due to the insufficient amount of generated asymmetry and the red region is excluded by BBN.

There are several things we would like to point out. First, the final baryon asymmetry only mildly depends on the involved particle masses and in general larger mass scales will only lead to a slight decrease of δY_B as can be seen from the slope of the BAU line in figure 9. In contrast, choosing a higher mass scale weakens the BBN bounds, e.g. for a DM mass of 6 keV we need $\Omega h_{N_2}^2 \lesssim 8.5$ for $\mu_2 = 600$ GeV, while for $\mu_2 = 800$ GeV the limit is relaxed to $\Omega h_{N_2}^2 \lesssim 88.7$. These values also suggest that the N_2 decay contribution ($\Omega h_{N_2 \rightarrow N_1}^2$) to the DM relic abundance is negligible and thus y_1 is fixed by the requirement that DM is completely produced via N_1 freeze-in. Larger scalar masses lead to stronger Yukawa interactions in order to maintain the correct relic density (see eq. (3.9)).

In figure 9, the blue region indicates the region excluded for $\delta_M \leq 10^{-11}$. A smaller level of degeneracy between right-handed neutrinos, such as $\delta_M = 10^{-10}$, would correspond to the shift of this exclusion to the left (toward BBN bound). For $\delta_M \geq 10^{-10}$ and $m_{N_1} \sim 6$ keV we find no parameter space free from either BBN or BAU exclusions.

We have also observed the dependence of BBN limits on the maximum allowed DM mass. For μ_2 in the range [300, 1000] GeV, we found a maximal value of the DM mass of 9.4 keV, which is consistent with structure formation bounds. Generally, choosing higher degeneracies will open up the available parameter space, thus allowing larger DM masses. However, even in such cases we estimated the maximal allowed DM mass to be at most $\mathcal{O}(10)$ keV.

In summary, we identified the parameter space in which the produced DM abundance and BAU are in accord with the observed values. We have seen that the most stringent constraints arise from BBN considerations, which can be relaxed by employing coannihilation processes between $N_{2,3}$ and Σ which effectively put an upper bound on the allowed mass for the DM. We found that the DM production in our model is mainly driven by the freeze-in.

In this section we have shown that the three biggest challenges of particle physics, pointed out in section 1, can be successfully and simultaneously solved within the considered model.

6 Detection prospects

In section 2.1 we discussed the limits from structure formation on keV-scale DM. Here, we wish to point out that these limits are relaxed in this model which opens up some parameter space for such DM candidate. Let us illustrate how DM can be made colder, i.e. less constrained from structure formation considerations. We discuss the most relevant case where the observable DM relic density is dominantly set through the freeze-in from the decays of heavy scalar particles (see section 3.1), as the strong production from N_2 decays is in tension with BBN limits. The production of DM occurs at the mass scale of decaying charged and neutral scalars. Due to the production at such high temperatures, cooling of DM particles is efficient. Namely, the effective temperature of the DM sector, when compared to photons, is reduced by the amount of entropy dilution factor which is ≈ 2.9 [24, 92].

Having the absence of X-ray signal, we reach the conclusion that the testability of our model is currently limited only to the searches at the LHC as well as the facilities probing lepton flavor violation processes. The limits coming from the latter are discussed in section 2.1 and consistently taken into account throughout the paper. In this section, therefore, we mainly comment on the LHC prospects, which were studied for this model in [93] where the dominant production of DM is assumed to be via freeze-in through scalar decays. Among others, the authors are considering the regime where scalar particles are heavier than all three generations of right-handed neutrinos, which matches our setup.

The answer to the question which LHC search has the strongest sensitivity depends on the mass difference between m_{N_2} and m_{N_3} . In section 4, we showed that baryogenesis can be achieved via generation of a lepton asymmetry, where the crucial ingredient was the approximate degeneracy between N_2 and N_3 states. In that case, the LHC searches are limited to $\sigma^\pm \rightarrow N_{2,3} l_\alpha^\pm$, i.e. channels including charged leptons in the final state. Charged scalars σ^\pm are produced *in pairs* either via gluon fusion or Drell-Yan processes [93], and therefore the expected signature consists of two prompt charged leptons and missing energy due to elusive right-handed neutrinos which are decaying to N_1 only after leaving the detector.

Let us note that with an assumption of a complementary baryogenesis mechanism that goes beyond our framework (for instance decays of very heavy singlets [94] or electroweak baryogenesis [95]) for which tiny δ_M is not required, the discovery potential at LHC increases. Namely, if we assume a mass gap between N_2 and N_3 to be $\gtrsim 10$ GeV there is another viable search in this model — displaced vertices [96]. Even though decays of heavier right-handed neutrinos into DM (N_1) do not happen within the detector, the decay

$N_3 \rightarrow N_2 l_\alpha^\pm l_\beta^\mp$ can be rapid enough. Then, the displaced leptons from the initial scalar decay and the subsequent N_3 decay may be observed. Limits for both leptons+ \cancel{E}_T and displaced vertices are presented in [93]. Let us emphasize that displaced vertex search can constrain $y_{3\alpha}$ and $y_{2\alpha}$ couplings up to the level of 10^{-4} , which is two orders of magnitude stronger with respect to the limits from LFV assuming no strong hierarchy between the entries of the Yukawa matrix. The existence of the upper limit is a consequence of the requirement for the resolution of displaced vertices. It is worth mentioning that the limits from LHC searches attenuate very quickly above EW scale. For more details we refer the reader to ref. [93].

Let us also remark that the presence of extra charged scalars implies the radiative contribution to the decays of SM Higgs particle into a pair of photons or a photon and a Z boson [97, 98]. Specifically, for the diphoton channel, in order to get constructive (destructive) interference between SM and new physics contribution, one requires negative (positive) Higgs portal couplings. Significant deviation from measured values [40] are achieved for rather large values of such couplings [99, 100] which are avoided in our analysis.

We finalize this section with a comment on the detection prospects for $N_{2,3}$ discovery at future lepton colliders as well as SHIP [101], that were proposed recently [80, 102, 103]. At lepton colliders, right-handed neutrinos are assumed to be generated in the process $e^+e^- \rightarrow N_{2,3} \nu_\alpha$ [104], for which the mixing between active and right-handed states is required. The mixing is necessary also for SHIP where right-handed neutrinos would get produced in the decays of heavy hadrons. Hence, due to the absence of the mixing induced by an exact \mathbb{Z}_2 parity symmetry, right-handed neutrinos in this model are not testable at these facilities.

7 Summary and conclusions

In this work we studied the generation of neutrino masses, dark matter and baryon asymmetry of the Universe within the so called Scotogenic model, where three right-handed neutrinos and an additional scalar doublet, all odd under a \mathbb{Z}_2 parity symmetry, are added to the SM. Active neutrino masses are obtained radiatively via loops involving new particles. We considered a mass spectrum where all scalar masses are at or below the TeV-scale. Furthermore, we invoked a hierarchy in the right-handed neutrino mass spectrum, choosing $m_{N_1} \approx \mathcal{O}(1)$ keV and $m_{N_{2,3}} \approx \mathcal{O}(100)$ GeV, where the lightest state, N_1 , is a keV-scale DM particle.

For DM production, we looked at two complementary contributions: first, we examined freeze-in of N_1 from the decays of new scalars. Second, we also took three-body decays of “frozen-out” N_2 into account. We were able to derive the correct DM density in a wide parameter region.

As the baryon asymmetry is concerned, we studied leptogenesis from the combination of right-handed neutrino oscillations and scalar decays. We showed that it is possible to derive a significant asymmetry in case of highly degenerated right-handed neutrino masses. We have established that it is crucial to set the mass of heavier two states to $\gtrsim \mathcal{O}(10^2)$ GeV in order to prevent late time washout.

Finally, in finding the joint parameter space for DM and BAU, we have shown that the BBN bound plays a major role and rules out a large portion of the parameter space, effectively forbidding DM mass to exceed ~ 10 keV. Nevertheless, by imposing coannihilla-

tions between fermions and scalars, we were successful in identifying the regions where both DM and BAU are produced in right amounts. Hence, in the considered model, we have simultaneously solved the three greatest mysteries in particle physics.

Acknowledgments

We thank Joachim Kopp for the collaboration in the initial stage of this project. We are greatly indebted to Daniele Teresi for cross checking our leptogenesis implementation in the initial phase of this work as well as helping us understand several key leptogenesis features, in particular the production from scalar decays. VB would also like to thank Mikhail Shaposhnikov and Kai Schmitz for very insightful and useful discussions on the presented model. Work in Mainz is supported by the Cluster of Excellence Precision Physics, Fundamental Interactions and Structure of Matter (PRISMA — EXC 1098).

A Phase space integration and reaction densities

Here we present the results for the reaction densities and washout terms, obtained by performing the phase space integration. In the following, we make the abbreviation $E_i \equiv E_i(\mathbf{p}_i)$ for SM lepton (E_L), right-handed neutrino (E_N) and scalar (E_Σ) energies. The right-handed neutrino and SM lepton momenta are denoted with k and p , respectively. We also use the approximation $m_{N_2} \simeq m_{N_3} \equiv m_N$. The phase space integral is evaluated as

$$\begin{aligned}
 \int d\Pi_{PS} &= \int \frac{d^3\mathbf{k}}{(2\pi)^3} \frac{1}{2E_N} \int \frac{d^3\mathbf{p}}{(2\pi)^3} \frac{1}{2E_L} \frac{2\pi}{2E_\Sigma} \delta(E_\Sigma - E_N - E_L) \\
 &= \int_0^\infty \frac{dk k^2}{(2\pi)^3} \int d\Omega_k \frac{1}{2E_N} \int_0^\infty \frac{dp p^2}{(2\pi)^3} \int d\Omega_p \frac{1}{2E_L} \frac{2\pi}{2E_\Sigma} \delta(E_\Sigma - E_N - E_L) \\
 &= \frac{1}{32\pi^3} \int_0^\infty \frac{dk k^2}{E_N} \int_0^\infty \frac{dp p^2}{E_L} \int_{-1}^1 d\cos\theta_{12} \frac{1}{E_\Sigma} \delta(E_\Sigma - E_N - E_L) \\
 &= \frac{1}{32\pi^3} \int_{m_N}^\infty dE_N \sqrt{E_N^2 - m_N^2} \int_{M_L}^\infty dE_L \sqrt{E_L^2 - M_L^2} \int_{-1}^1 d\cos\theta_{12} \frac{1}{E_\Sigma} \\
 &\quad \times \delta\left(\sqrt{q^2 + (m_\pm)^2} - \sqrt{k^2 + m_N^2} - \sqrt{p^2 + M_L^2}\right) \tag{A.1} \\
 &\equiv \frac{1}{32\pi^3} \int_{m_N}^\infty dE_N \sqrt{E_N^2 - m_N^2} \int_{M_L}^\infty dE_L \sqrt{E_L^2 - M_L^2} \int_{-1}^1 d\cos\theta_{12} \frac{1}{E_\Sigma} \delta(f(\cos\theta_{12})),
 \end{aligned}$$

where θ_{12} is the angle between the outgoing lepton and right-handed neutrino and

$$f(\cos\theta_{12}) \equiv \sqrt{p^2 + k^2 + 2pk \cos\theta_{12} + m_\pm^2} - \sqrt{k^2 + m_N^2} - \sqrt{p^2 + M_L^2}. \tag{A.2}$$

After a further evaluation we reach the expression

$$\int d\Pi_{PS} = \frac{1}{32\pi^3} \int_{m_N}^{\infty} dE_N \int_{E^-}^{E^+} dE_L \int d\cos\theta_{12} \delta(\cos\theta_{12} - \cos\theta_{12}^{\min}), \quad (\text{A.3})$$

with

$$E^{\pm} \equiv \pm \frac{\sqrt{\sqrt{E_N^2 - m_N^2} (m_{\pm}^2 - (M_L + m_N)^2) (m_{\pm}^2 - (M_L - m_N)^2)}}{2m_N^2} \mp \frac{E_N (m_N^2 + M_L^2 - m_{\pm}^2)}{2m_N^2}, \quad (\text{A.4})$$

and

$$\cos\theta_{12}^{\min} \equiv \frac{2E_L E_N + m_N^2 + M_L^2 - m_{\pm}^2}{2k\sqrt{E_N^2 - m_N^2}}. \quad (\text{A.5})$$

All reaction densities and washout terms include

$$\begin{aligned} & \frac{2}{32\pi^3} \int_{M_N}^{\infty} dE_N \int_{E^-}^{E^+} dE_L \left(\frac{E_N}{2\sqrt{E_N^2 - M_N^2}} \pm \frac{1}{2} \right) \\ & \times \left[2\sqrt{E_N^2 - M_N^2} E_L \mp (2E_N E_L + M_N^2 + M_L^2 - m_{\pm}^2) \right]. \end{aligned} \quad (\text{A.6})$$

After inserting appropriate distribution functions we finally obtain

$$\begin{aligned} \gamma_{\alpha,ij}^{LC} &= \frac{2}{32\pi^3} \int_{m_N}^{\infty} dE_N \int_{E^-}^{E^+} dE_L \frac{1}{e^{E_N/T} + 1} \left(\frac{1}{e^{E_L/T} + 1} + \frac{1}{e^{(E_L+E_N)/T} - 1} \right) \\ & \times \left(\frac{E_N}{2\sqrt{E_N^2 - m_N^2}} + \frac{1}{2} \right) \left[2\sqrt{E_N^2 - m_N^2} E_L - (2E_N E_L + m_N^2 + M_L^2 - m_{\pm}^2) \right] y_{\alpha i}^T y_{\alpha j}^{\dagger}. \end{aligned} \quad (\text{A.7})$$

$$\begin{aligned} \gamma_{\alpha,ij}^{LV} &= \frac{2}{32\pi^3} \int_{m_N}^{\infty} dE_N \int_{E^-}^{E^+} dE_L \frac{1}{e^{E_N/T} + 1} \left(\frac{1}{e^{E_L/T} + 1} + \frac{1}{e^{(E_L+E_N)/T} - 1} \right) \\ & \times \left(\frac{E_N}{2\sqrt{E_N^2 - m_N^2}} - \frac{1}{2} \right) \left[2\sqrt{E_N^2 - m_N^2} E_L + (2E_N E_L + m_N^2 + M_L^2 - m_{\pm}^2) \right] y_{\alpha j}^T y_{\alpha i}^{\dagger}. \end{aligned} \quad (\text{A.8})$$

$$\begin{aligned} \gamma_{WQ,\alpha,ij}^{LC} &= \frac{2}{32\pi^3} \int_{m_N}^{\infty} dE_N \int_{E^-}^{E^+} dE_L \left(\frac{1}{e^{E_L/T} + 1} \cdot \frac{1}{e^{(E_L+E_N)/T} - 1} \right) \\ & \times \left(\frac{E_N}{2\sqrt{E_N^2 - m_N^2}} + \frac{1}{2} \right) \left[2\sqrt{E_N^2 - m_N^2} E_L - (2E_N E_L + m_N^2 + M_L^2 - m_{\pm}^2) \right] y_{\alpha i}^T y_{\alpha j}^{\dagger}. \end{aligned} \quad (\text{A.9})$$

$$\gamma_{WQ,\alpha,ij}^{LV} = \frac{2}{32\pi^3} \int_{m_N}^{\infty} dE_N \int_{E^-}^{E^+} dE_L \left(\frac{1}{e^{E_L/T} + 1} \cdot \frac{1}{e^{(E_L+E_N)/T} - 1} \right) \quad (\text{A.10})$$

$$\times \left(\frac{E_N}{2\sqrt{E_N^2 - m_N^2}} - \frac{1}{2} \right) \left[2\sqrt{E_N^2 - m_N^2} E_L + (2E_N E_L + m_N^2 + M_L^2 - m_{\pm}^2) \right] y_{\alpha j}^T y_{\alpha i}^{\dagger}.$$

$$\gamma_{WC,\alpha,ij}^{LC} = \frac{2}{32\pi^3} \int_{m_N}^{\infty} dE_N \int_{E^-}^{E^+} dE_L \left(\frac{1}{e^{E_N/T} + 1} \cdot \frac{1}{e^{E_L/T} - 1} \right) \quad (\text{A.11})$$

$$\times \left(\frac{E_N}{2\sqrt{E_N^2 - m_N^2}} + \frac{1}{2} \right) \left[2\sqrt{E_N^2 - m_N^2} E_L - (2E_N E_L + m_N^2 + M_L^2 - m_{\pm}^2) \right] y_{\alpha i}^T y_{\alpha j}^{\dagger}.$$

$$\gamma_{WC,\alpha,ij}^{LV} = \frac{2}{32\pi^3} \int_{m_N}^{\infty} dE_N \int_{E^-}^{E^+} dE_L \left(\frac{1}{e^{E_N/T} + 1} \cdot \frac{1}{e^{E_L/T} - 1} \right) \quad (\text{A.12})$$

$$\times \left(\frac{E_N}{2\sqrt{E_N^2 - m_N^2}} - \frac{1}{2} \right) \left[2\sqrt{E_N^2 - m_N^2} E_L + (2E_N E_L + m_N^2 + M_L^2 - m_{\pm}^2) \right] y_{\alpha j}^T y_{\alpha i}^{\dagger}.$$

Open Access. This article is distributed under the terms of the Creative Commons Attribution License ([CC-BY 4.0](https://creativecommons.org/licenses/by/4.0/)), which permits any use, distribution and reproduction in any medium, provided the original author(s) and source are credited.

References

- [1] E. Ma, *Verifiable radiative seesaw mechanism of neutrino mass and dark matter*, *Phys. Rev. D* **73** (2006) 077301 [[hep-ph/0601225](#)] [[INSPIRE](#)].
- [2] E.K. Akhmedov, V.A. Rubakov and A. Yu. Smirnov, *Baryogenesis via neutrino oscillations*, *Phys. Rev. Lett.* **81** (1998) 1359 [[hep-ph/9803255](#)] [[INSPIRE](#)].
- [3] T. Hambye and D. Teresi, *Higgs doublet decay as the origin of the baryon asymmetry*, *Phys. Rev. Lett.* **117** (2016) 091801 [[arXiv:1606.00017](#)] [[INSPIRE](#)].
- [4] T. Asaka, S. Blanchet and M. Shaposhnikov, *The ν MSM, dark matter and neutrino masses*, *Phys. Lett. B* **631** (2005) 151 [[hep-ph/0503065](#)] [[INSPIRE](#)].
- [5] T. Asaka and M. Shaposhnikov, *The ν MSM, dark matter and baryon asymmetry of the universe*, *Phys. Lett. B* **620** (2005) 17 [[hep-ph/0505013](#)] [[INSPIRE](#)].
- [6] M. Shaposhnikov, *The ν MSM, leptonic asymmetries and properties of singlet fermions*, *JHEP* **08** (2008) 008 [[arXiv:0804.4542](#)] [[INSPIRE](#)].
- [7] X.-D. Shi and G.M. Fuller, *A new dark matter candidate: nonthermal sterile neutrinos*, *Phys. Rev. Lett.* **82** (1999) 2832 [[astro-ph/9810076](#)] [[INSPIRE](#)].
- [8] S. Dodelson and L.M. Widrow, *Sterile-neutrinos as dark matter*, *Phys. Rev. Lett.* **72** (1994) 17 [[hep-ph/9303287](#)] [[INSPIRE](#)].
- [9] R.N. Mohapatra and G. Senjanović, *Neutrino mass and spontaneous parity violation*, *Phys. Rev. Lett.* **44** (1980) 912 [[INSPIRE](#)].

- [10] P. Minkowski, $\mu \rightarrow e\gamma$ at a rate of one out of 10^9 muon decays?, *Phys. Lett. B* **67** (1977) 421 [INSPIRE].
- [11] M. Gell-Mann, P. Ramond and R. Slansky, *Complex spinors and unified theories*, *Conf. Proc. C* **790927** (1979) 315 [arXiv:1306.4669] [INSPIRE].
- [12] T. Yanagida, *Horizontal symmetry and masses of neutrinos*, *Conf. Proc. C* **7902131** (1979) 95 [INSPIRE].
- [13] J. Baur et al., *Constraints from Ly- α forests on non-thermal dark matter including resonantly-produced sterile neutrinos*, *JCAP* **12** (2017) 013 [arXiv:1706.03118] [INSPIRE].
- [14] K. Perez, K.C.Y. Ng, J.F. Beacom, C. Hersh, S. Horiuchi and R. Krivonos, *Almost closing the ν MSM sterile neutrino dark matter window with NuSTAR*, *Phys. Rev. D* **95** (2017) 123002 [arXiv:1609.00667] [INSPIRE].
- [15] A. Merle, A. Schneider and M. Totzauer, *Dodelson-Widrow production of sterile neutrino dark matter with non-trivial initial abundance*, *JCAP* **04** (2016) 003 [arXiv:1512.05369] [INSPIRE].
- [16] S. Horiuchi, P.J. Humphrey, J. Onorbe, K.N. Abazajian, M. Kaplinghat and S. Garrison-Kimmel, *Sterile neutrino dark matter bounds from galaxies of the local group*, *Phys. Rev. D* **89** (2014) 025017 [arXiv:1311.0282] [INSPIRE].
- [17] D. Malyshev, A. Neronov and D. Eckert, *Constraints on 3.55 keV line emission from stacked observations of dwarf spheroidal galaxies*, *Phys. Rev. D* **90** (2014) 103506 [arXiv:1408.3531] [INSPIRE].
- [18] T. Tamura, R. Iizuka, Y. Maeda, K. Mitsuda and N.Y. Yamasaki, *An X-ray spectroscopic search for dark matter in the Perseus cluster with Suzaku*, *Publ. Astron. Soc. Jap.* **67** (2015) 23 [arXiv:1412.1869] [INSPIRE].
- [19] HITOMI collaboration, F.A. Aharonian et al., *Hitomi constraints on the 3.5 keV line in the Perseus galaxy cluster*, *Astrophys. J.* **837** (2017) L15 [arXiv:1607.07420] [INSPIRE].
- [20] E. Bulbul, M. Markevitch, A. Foster, R.K. Smith, M. Loewenstein and S.W. Randall, *Detection of an unidentified emission line in the stacked X-ray spectrum of galaxy clusters*, *Astrophys. J.* **789** (2014) 13 [arXiv:1402.2301] [INSPIRE].
- [21] A. Boyarsky, O. Ruchayskiy, D. Iakubovskiy and J. Franse, *Unidentified line in X-ray spectra of the Andromeda galaxy and Perseus galaxy cluster*, *Phys. Rev. Lett.* **113** (2014) 251301 [arXiv:1402.4119] [INSPIRE].
- [22] N. Cappelluti et al., *Searching for the 3.5 keV line in the deep fields with Chandra: the 10 Ms observations*, *Astrophys. J.* **854** (2018) 179 [arXiv:1701.07932] [INSPIRE].
- [23] K.N. Abazajian, *Sterile neutrinos in cosmology*, *Phys. Rept.* **711-712** (2017) 1 [arXiv:1705.01837] [INSPIRE].
- [24] V. Brdar, J. Kopp, J. Liu and X.-P. Wang, *X-ray lines from dark matter annihilation at the keV scale*, *Phys. Rev. Lett.* **120** (2018) 061301 [arXiv:1710.02146] [INSPIRE].
- [25] T.E. Jeltema and S. Profumo, *Discovery of a 3.5 keV line in the galactic centre and a critical look at the origin of the line across astronomical targets*, *Mon. Not. Roy. Astron. Soc.* **450** (2015) 2143 [arXiv:1408.1699] [INSPIRE].

- [26] T. Jeltema and S. Profumo, *Reply to two comments on “dark matter searches going bananas the contribution of potassium (and chlorine) to the 3.5 keV line”*, [arXiv:1411.1759](#) [INSPIRE].
- [27] C. Shah et al., *Laboratory measurements compellingly support charge-exchange mechanism for the ‘dark matter’ ~ 3.5 keV X-ray line*, *Astrophys. J.* **833** (2016) 52 [[arXiv:1608.04751](#)] [INSPIRE].
- [28] R. Barbieri, L.J. Hall and V.S. Rychkov, *Improved naturalness with a heavy Higgs: an alternative road to LHC physics*, *Phys. Rev. D* **74** (2006) 015007 [[hep-ph/0603188](#)] [INSPIRE].
- [29] G.C. Branco, P.M. Ferreira, L. Lavoura, M.N. Rebelo, M. Sher and J.P. Silva, *Theory and phenomenology of two-Higgs-doublet models*, *Phys. Rept.* **516** (2012) 1 [[arXiv:1106.0034](#)] [INSPIRE].
- [30] D. Tytler, J.M. O’Meara, N. Suzuki and D. Lubin, *Review of big bang nucleosynthesis and primordial abundances*, *Phys. Scripta T* **85** (2000) 12 [[astro-ph/0001318](#)] [INSPIRE].
- [31] M. Pettini and R. Cooke, *A new, precise measurement of the primordial abundance of deuterium*, *Mon. Not. Roy. Astron. Soc.* **425** (2012) 2477 [[arXiv:1205.3785](#)] [INSPIRE].
- [32] M. Kawasaki, K. Kohri, T. Moroi and Y. Takaesu, *Revisiting big-bang nucleosynthesis constraints on long-lived decaying particles*, *Phys. Rev. D* **97** (2018) 023502 [[arXiv:1709.01211](#)] [INSPIRE].
- [33] E. Molinaro, C.E. Yaguna and O. Zapata, *FIMP realization of the scotogenic model*, *JCAP* **07** (2014) 015 [[arXiv:1405.1259](#)] [INSPIRE].
- [34] A. Schneider, *Astrophysical constraints on resonantly produced sterile neutrino dark matter*, *JCAP* **04** (2016) 059 [[arXiv:1601.07553](#)] [INSPIRE].
- [35] A. Schneider, *Structure formation with suppressed small-scale perturbations*, *Mon. Not. Roy. Astron. Soc.* **451** (2015) 3117 [[arXiv:1412.2133](#)] [INSPIRE].
- [36] A. Schneider, *Testing the sterile neutrino dark matter paradigm with astrophysical observations*, *PoS(NOW2016)093* [[arXiv:1704.01832](#)] [INSPIRE].
- [37] J.F. Cherry and S. Horiuchi, *Closing in on resonantly produced sterile neutrino dark matter*, *Phys. Rev. D* **95** (2017) 083015 [[arXiv:1701.07874](#)] [INSPIRE].
- [38] G. Kulkarni, J.F. Hennawi, J. Oñorbe, A. Rorai and V. Springel, *Characterizing the pressure smoothing scale of the intergalactic medium*, *Astrophys. J.* **812** (2015) 30 [[arXiv:1504.00366](#)] [INSPIRE].
- [39] T. Toma and A. Vicente, *Lepton flavor violation in the scotogenic model*, *JHEP* **01** (2014) 160 [[arXiv:1312.2840](#)] [INSPIRE].
- [40] PARTICLE DATA GROUP collaboration, C. Patrignani et al., *Review of particle physics*, *Chin. Phys. C* **40** (2016) 100001 [INSPIRE].
- [41] M. Lindner, M. Platscher, C.E. Yaguna and A. Merle, *Fermionic WIMPs and vacuum stability in the scotogenic model*, *Phys. Rev. D* **94** (2016) 115027 [[arXiv:1608.00577](#)] [INSPIRE].
- [42] A. Ahriche, A. Jueid and S. Nasri, *Radiative neutrino mass and Majorana dark matter within an inert Higgs doublet model*, *Phys. Rev. D* **97** (2018) 095012 [[arXiv:1710.03824](#)] [INSPIRE].

- [43] M.E. Peskin and T. Takeuchi, *Estimation of oblique electroweak corrections*, *Phys. Rev. D* **46** (1992) 381 [INSPIRE].
- [44] A. Zee, *A theory of lepton number violation, neutrino Majorana mass and oscillation*, *Phys. Lett. B* **93** (1980) 389 [Erratum *ibid.* **B 95** (1980) 461] [INSPIRE].
- [45] A. Zee, *Charged scalar field and quantum number violations*, *Phys. Lett. B* **161** (1985) 141 [INSPIRE].
- [46] A. Zee, *Quantum numbers of Majorana neutrino masses*, *Nucl. Phys. B* **264** (1986) 99 [INSPIRE].
- [47] K.S. Babu and V.S. Mathur, *Magnetic moments of Dirac and Majorana neutrinos*, *Phys. Lett. B* **196** (1987) 218 [INSPIRE].
- [48] V. Brdar, I. Picek and B. Radovic, *Radiative neutrino mass with scotogenic scalar triplet*, *Phys. Lett. B* **728** (2014) 198 [arXiv:1310.3183] [INSPIRE].
- [49] K. Kumericki, T. Mede and I. Picek, *Renormalizable SU(5) completions of a Zee-type neutrino mass model*, *Phys. Rev. D* **97** (2018) 055012 [arXiv:1712.05246] [INSPIRE].
- [50] A. Das, T. Nomura, H. Okada and S. Roy, *Generation of a radiative neutrino mass in the linear seesaw framework, charged lepton flavor violation and dark matter*, *Phys. Rev. D* **96** (2017) 075001 [arXiv:1704.02078] [INSPIRE].
- [51] T. Appelquist and R. Shrock, *Neutrino masses in theories with dynamical electroweak symmetry breaking*, *Phys. Lett. B* **548** (2002) 204 [hep-ph/0204141] [INSPIRE].
- [52] T. Appelquist and R. Shrock, *Dynamical symmetry breaking of extended gauge symmetries*, *Phys. Rev. Lett.* **90** (2003) 201801 [hep-ph/0301108] [INSPIRE].
- [53] SUPER-KAMIOKANDE collaboration, Y. Fukuda et al., *Evidence for oscillation of atmospheric neutrinos*, *Phys. Rev. Lett.* **81** (1998) 1562 [hep-ex/9807003] [INSPIRE].
- [54] T. Kajita, *Nobel lecture: discovery of atmospheric neutrino oscillations*, *Rev. Mod. Phys.* **88** (2016) 030501 [INSPIRE].
- [55] J.A. Casas and A. Ibarra, *Oscillating neutrinos and $\mu \rightarrow e, \gamma$* , *Nucl. Phys. B* **618** (2001) 171 [hep-ph/0103065] [INSPIRE].
- [56] PARTICLE DATA GROUP collaboration, C. Patrignani et al., *Review of particle physics*, *Chin. Phys. C* **40** (2016) 100001 [INSPIRE].
- [57] L.J. Hall, K. Jedamzik, J. March-Russell and S.M. West, *Freeze-in production of FIMP dark matter*, *JHEP* **03** (2010) 080 [arXiv:0911.1120] [INSPIRE].
- [58] L. Calibbi, L. Lopez-Honorez, S. Lowette and A. Mariotti, *Singlet-doublet dark matter freeze-in: LHC displaced signatures versus cosmology*, arXiv:1805.04423 [INSPIRE].
- [59] A. Biswas, D. Borah and A. Dasgupta, *A UV complete framework of freeze-in massive particle dark matter*, arXiv:1805.06903 [INSPIRE].
- [60] M. Lisanti, *Lectures on dark matter physics*, in *Proceedings, Theoretical Advanced Study Institute in Elementary Particle Physics: new frontiers in fields and strings (TASI 2015)*, Boulder, CO, U.S.A., 1–26 June 2015, *World Scientific*, Singapore, (2017), pg. 399 [arXiv:1603.03797] [INSPIRE].
- [61] F. Bezrukov, H. Hettmansperger and M. Lindner, *keV sterile neutrino dark matter in gauge extensions of the Standard Model*, *Phys. Rev. D* **81** (2010) 085032 [arXiv:0912.4415] [INSPIRE].

- [62] S.F. King and A. Merle, *Warm dark matter from keVins*, *JCAP* **08** (2012) 016 [[arXiv:1205.0551](#)] [[INSPIRE](#)].
- [63] A. Merle and M. Platscher, *Parity problem of the scotogenic neutrino model*, *Phys. Rev. D* **92** (2015) 095002 [[arXiv:1502.03098](#)] [[INSPIRE](#)].
- [64] K.N. Abazajian, *Detection of dark matter decay in the X-ray*, in *Astro2010: the astronomy and astrophysics decadal survey*, National Research Council, U.S.A., (2010) [[arXiv:0903.2040](#)] [[INSPIRE](#)].
- [65] J. Edsjo and P. Gondolo, *Neutralino relic density including coannihilations*, *Phys. Rev. D* **56** (1997) 1879 [[hep-ph/9704361](#)] [[INSPIRE](#)].
- [66] G. Bélanger, F. Boudjema and A. Pukhov, *MicrOMEGAs: a code for the calculation of dark matter properties in generic models of particle interaction*, in *The Dark Secrets of the Terascale: Proceedings, TASI 2011*, Boulder, CO, U.S.A., 6 June–11 July 2011, *World Scientific*, Singapore, (2013), pg. 739 [[arXiv:1402.0787](#)] [[INSPIRE](#)].
- [67] D. Barducci et al., *Collider limits on new physics within MicrOMEGAs 4.3*, *Comput. Phys. Commun.* **222** (2018) 327 [[arXiv:1606.03834](#)] [[INSPIRE](#)].
- [68] M. Backovic, K. Kong and M. McCaskey, *MadDM v.1.0: computation of dark matter relic abundance using MadGraph5*, *Physics of the Dark Universe* **5-6** (2014) 18 [[arXiv:1308.4955](#)] [[INSPIRE](#)].
- [69] PLANCK collaboration, P.A.R. Ade et al., *Planck 2015 results. XIII. Cosmological parameters*, *Astron. Astrophys.* **594** (2016) A13 [[arXiv:1502.01589](#)] [[INSPIRE](#)].
- [70] W. Buchmüller, P. Di Bari and M. Plümacher, *Leptogenesis for pedestrians*, *Annals Phys.* **315** (2005) 305 [[hep-ph/0401240](#)] [[INSPIRE](#)].
- [71] V.A. Kuzmin, V.A. Rubakov and M.E. Shaposhnikov, *On the anomalous electroweak baryon number nonconservation in the early universe*, *Phys. Lett. B* **155** (1985) 36 [[INSPIRE](#)].
- [72] M. D’Onofrio, K. Rummukainen and A. Tranberg, *Sphaleron rate in the minimal Standard Model*, *Phys. Rev. Lett.* **113** (2014) 141602 [[arXiv:1404.3565](#)] [[INSPIRE](#)].
- [73] W. Buchmüller, P. Di Bari and M. Plümacher, *Cosmic microwave background, matter-antimatter asymmetry and neutrino masses*, *Nucl. Phys. B* **643** (2002) 367 [*Erratum ibid.* **B 793** (2008) 362] [[hep-ph/0205349](#)] [[INSPIRE](#)].
- [74] J.D. Clarke, R. Foot and R.R. Volkas, *Electroweak naturalness in the three-flavor type-I seesaw model and implications for leptogenesis*, *Phys. Rev. D* **91** (2015) 073009 [[arXiv:1502.01352](#)] [[INSPIRE](#)].
- [75] R. Barbieri, P. Creminelli, A. Strumia and N. Tetradis, *Baryogenesis through leptogenesis*, *Nucl. Phys. B* **575** (2000) 61 [[hep-ph/9911315](#)] [[INSPIRE](#)].
- [76] A. Abada, S. Davidson, F.-X. Josse-Michaux, M. Losada and A. Riotto, *Flavor issues in leptogenesis*, *JCAP* **04** (2006) 004 [[hep-ph/0601083](#)] [[INSPIRE](#)].
- [77] A. De Simone and A. Riotto, *On the impact of flavour oscillations in leptogenesis*, *JCAP* **02** (2007) 005 [[hep-ph/0611357](#)] [[INSPIRE](#)].
- [78] M. Beneke, B. Garbrecht, C. Fidler, M. Herranen and P. Schwaller, *Flavoured leptogenesis in the CTP formalism*, *Nucl. Phys. B* **843** (2011) 177 [[arXiv:1007.4783](#)] [[INSPIRE](#)].

- [79] K. Moffat, S. Pascoli, S.T. Petcov, H. Schulz and J. Turner, *Three-flavored nonresonant leptogenesis at intermediate scales*, *Phys. Rev. D* **98** (2018) 015036 [[arXiv:1804.05066](#)] [[INSPIRE](#)].
- [80] P. Hernández, M. Kekic, J. López-Pavón, J. Racker and J. Salvado, *Testable baryogenesis in seesaw models*, *JHEP* **08** (2016) 157 [[arXiv:1606.06719](#)] [[INSPIRE](#)].
- [81] M. Drewes, B. Garbrecht, D. Gueter and J. Klaric, *Leptogenesis from oscillations of heavy neutrinos with large mixing angles*, *JHEP* **12** (2016) 150 [[arXiv:1606.06690](#)] [[INSPIRE](#)].
- [82] M. Drewes et al., *ARS leptogenesis*, *Int. J. Mod. Phys. A* **33** (2018) 1842002 [[arXiv:1711.02862](#)] [[INSPIRE](#)].
- [83] B. Shuve and I. Yavin, *Baryogenesis through neutrino oscillations: a unified perspective*, *Phys. Rev. D* **89** (2014) 075014 [[arXiv:1401.2459](#)] [[INSPIRE](#)].
- [84] T. Hambye and D. Teresi, *Baryogenesis from L-violating Higgs-doublet decay in the density-matrix formalism*, *Phys. Rev. D* **96** (2017) 015031 [[arXiv:1705.00016](#)] [[INSPIRE](#)].
- [85] D. Suematsu, *Thermal leptogenesis in a TeV scale model for neutrino masses*, *Eur. Phys. J. C* **72** (2012) 1951 [[arXiv:1103.0857](#)] [[INSPIRE](#)].
- [86] S. Kashiwase and D. Suematsu, *Baryon number asymmetry and dark matter in the neutrino mass model with an inert doublet*, *Phys. Rev. D* **86** (2012) 053001 [[arXiv:1207.2594](#)] [[INSPIRE](#)].
- [87] J. Racker, *Mass bounds for baryogenesis from particle decays and the inert doublet model*, *JCAP* **03** (2014) 025 [[arXiv:1308.1840](#)] [[INSPIRE](#)].
- [88] T. Hugle, M. Platscher and K. Schmitz, *Low-scale leptogenesis in the scotogenic neutrino mass model*, *Phys. Rev. D* **98** (2018) 023020 [[arXiv:1804.09660](#)] [[INSPIRE](#)].
- [89] P.S. Bhupal Dev, P. Millington, A. Pilaftsis and D. Teresi, *Flavour covariant transport equations: an application to resonant leptogenesis*, *Nucl. Phys. B* **886** (2014) 569 [[arXiv:1404.1003](#)] [[INSPIRE](#)].
- [90] G. Gil, P. Chankowski and M. Krawczyk, *Inert dark matter and strong electroweak phase transition*, *Phys. Lett. B* **717** (2012) 396 [[arXiv:1207.0084](#)] [[INSPIRE](#)].
- [91] N. Blinov, S. Profumo and T. Stefaniak, *The electroweak phase transition in the inert doublet model*, *JCAP* **07** (2015) 028 [[arXiv:1504.05949](#)] [[INSPIRE](#)].
- [92] J. Heeck and D. Teresi, *Cold keV dark matter from decays and scatterings*, *Phys. Rev. D* **96** (2017) 035018 [[arXiv:1706.09909](#)] [[INSPIRE](#)].
- [93] A.G. Hessler, A. Ibarra, E. Molinaro and S. Vogl, *Probing the scotogenic FIMP at the LHC*, *JHEP* **01** (2017) 100 [[arXiv:1611.09540](#)] [[INSPIRE](#)].
- [94] M. Fukugita and T. Yanagida, *Baryogenesis without grand unification*, *Phys. Lett. B* **174** (1986) 45 [[INSPIRE](#)].
- [95] D.E. Morrissey and M.J. Ramsey-Musolf, *Electroweak baryogenesis*, *New J. Phys.* **14** (2012) 125003 [[arXiv:1206.2942](#)] [[INSPIRE](#)].
- [96] O. Buchmueller et al., *Simplified models for displaced dark matter signatures*, *JHEP* **09** (2017) 076 [[arXiv:1704.06515](#)] [[INSPIRE](#)].
- [97] B. Swiezewska and M. Krawczyk, *Diphoton rate in the inert doublet model with a 125 GeV Higgs boson*, *Phys. Rev. D* **88** (2013) 035019 [[arXiv:1212.4100](#)] [[INSPIRE](#)].

- [98] A. Arhrib, R. Benbrik and N. Gaur, $H \rightarrow \gamma\gamma$ in inert Higgs doublet model, *Phys. Rev. D* **85** (2012) 095021 [[arXiv:1201.2644](#)] [[INSPIRE](#)].
- [99] M. Carena, I. Low and C.E.M. Wagner, Implications of a modified Higgs to diphoton decay width, *JHEP* **08** (2012) 060 [[arXiv:1206.1082](#)] [[INSPIRE](#)].
- [100] I. Picek and B. Radovic, Enhancement of $h \rightarrow \gamma\gamma$ by seesaw-motivated exotic scalars, *Phys. Lett. B* **719** (2013) 404 [[arXiv:1210.6449](#)] [[INSPIRE](#)].
- [101] O. Lantwin, Search for new physics with the SHiP experiment at CERN, *PoS(EPS-HEP2017)304* [[arXiv:1710.03277](#)] [[INSPIRE](#)].
- [102] FCC-EE STUDY TEAM collaboration, A. Blondel, E. Graverini, N. Serra and M. Shaposhnikov, Search for heavy right handed neutrinos at the FCC-ee, *Nucl. Part. Phys. Proc.* **273-275** (2016) 1883 [[arXiv:1411.5230](#)] [[INSPIRE](#)].
- [103] S. Antusch et al., Probing leptogenesis at future colliders, [arXiv:1710.03744](#) [[INSPIRE](#)].
- [104] S. Banerjee, P.S.B. Dev, A. Ibarra, T. Mandal and M. Mitra, Prospects of heavy neutrino searches at future lepton colliders, *Phys. Rev. D* **92** (2015) 075002 [[arXiv:1503.05491](#)] [[INSPIRE](#)].