Correction to: Observations of a Radio-Quiet Solar Preflare

Arnold O. Benz^{1,2} · Marina Battaglia¹ ·
Manuel Güdel³

Published online: 5 February 2018

© Springer Science+Business Media B.V., part of Springer Nature 2018

Correction to: Solar Phys

<https://doi.org/10.1007/s11207-017-1175-3>

This article was published with an erroneous version of Figure 3. Please find in this document the correct version of the figure.

Combined Radio and Space-based Solar Observations: From Techniques to New Results
Guest Editors: Eduard Kontar and Alexander Nindos

The online version of the original article can be found under <https://doi.org/10.1007/s11207-017-1175-3>.

✉ A.O. Benz
benz@astro.phys.ethz.ch

M. Battaglia
marina.battaglia@fhnw.ch

M. Güdel
manuel.guedel@univie.ac.at

¹ Institute for 4D Technologies, FHNW, 5210 Windisch, Switzerland

² Institute for Astronomy, ETH Zurich, 8093 Zürich, Switzerland

³ Department of Astrophysics, University of Vienna, Türkenschanzstrasse 17, 1180 Vienna, Austria

Figure 3 Comparison of the peak soft X-ray and 3.6 cm radio luminosity of solar and stellar flares. The upper limits derived from these preflare observations are shown with *red arrows*. Two peaks in the early main flare phase are indicated with a *lowercase red x*, the main peak with an *uppercase X*. We also show nanoflares in quiet regions (Krucker and Benz, 2000), microflares in active regions (Benz *et al.*, 1981; Fürst, Benz, and Hirth, 1982; Gary, Hartl, and Shimizu, 1997), impulsive and gradual flares (Benz and Güdel, 1994), quiescent BY Dra, dMe, and dKe stars (Güdel *et al.*, 1993), and a long-duration flare on the dMe star EQ Peg (Kundu *et al.*, 1988). The *dashed line* presents the relation given in Equation 1. (Graphics adapted from Benz and Güdel, 2010, and based on Krucker and Benz, 2000.)